

Deep 1.1 mm-wavelength imaging of the GOODS-S field by AzTEC/ASTE – I. Source catalogue and number counts

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ABSTRACT

We present the first results from a 1.1 mm confusion-limited map of the Great Observatories Origins Deep Survey-South (GOODS-S) taken with the AzTEC camera on the Atacama Submillimeter Telescope Experiment. We imaged a 270 arcmin^2 field to a 1σ depth of $0.48\text{--}0.73 \text{ mJy beam}^{-1}$, making this one of the deepest blank-field surveys at mm-wavelengths ever achieved. Although by traditional standards our GOODS-S map is extremely confused due to a sea of faint underlying sources, we demonstrate through simulations that our source identification and number counts analyses are robust, and the techniques discussed in this paper are relevant for other deeply confused surveys. We find a total of 41 dusty starburst galaxies with signal-to-noise ratios $S/N \geq 3.5$ within this uniformly covered region, where only two are expected to be false detections, and an additional seven robust source candidates located in the noisier ($1\sigma \approx 1 \text{ mJy beam}^{-1}$) outer region of the map. We derive the 1.1 mm number counts from this field using two different methods: a fluctuation or “ $P(d)$ ” analysis and a semi-Bayesian technique and find that both methods give consistent results. Our data are well fit by a Schechter function model with $(S', N_{3 \text{ mJy}}, \alpha) = (1.30^{+0.19}_{-0.25} \text{ mJy}, 160^{+27}_{-28} \text{ mJy}^{-1} \text{ deg}^{-2}, -2.0)$. Given the depth of this survey, we put the first tight constraints on the 1.1 mm number counts at $S_{1.1 \text{ mm}} = 0.5 \text{ mJy}$, and we find evidence that the faint end of the number counts at $S_{850 \text{ }\mu\text{m}} \lesssim 2.0 \text{ mJy}$ from various SCUBA surveys towards lensing clusters are biased high. In contrast to the $870 \text{ }\mu\text{m}$ survey of this field with the LABOCA camera, we find no apparent underdensity of sources compared to previous surveys at 1.1 mm; the estimates of the number counts of SMGs at flux densities $> 1 \text{ mJy}$ determined here are consistent with those measured

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from the AzTEC/SHADES survey. Additionally, we find a significant number of SMGs not identified in the LABOCA catalogue. We find that in contrast to observations at $\lambda \leq 500 \mu\text{m}$, MIPS 24 μm sources do not resolve the total energy density in the cosmic infrared background at 1.1 mm, demonstrating that a population of $z \gtrsim 3$ dust-obscured galaxies that are unaccounted for at these shorter wavelengths potentially contribute to a large fraction ($\sim 2/3$) of the infrared background at 1.1 mm.

Key words: methods: data analysis – surveys – galaxies: evolution – galaxies: high-redshift – galaxies: starburst – submillimetre: galaxies.

1 INTRODUCTION

Galaxies selected at submillimetre (sub-mm) and millimetre (mm) wavelengths (hereafter SMGs) comprise a population of dust-obscured starburst or active galactic nuclei (AGN) host galaxies at high redshift ($z \gtrsim 1$; see review by Blain et al. 2002). With far-infrared (FIR) luminosities $L_{\text{FIR}} \gtrsim 10^{12} L_{\odot}$, these systems appear to be scaled-up analogues to the ultraluminous infrared galaxies (ULIRGs) observed in the local Universe (Sanders & Mirabel 1996). Their FIR to mm spectral energy distributions (SEDs) are characterized by thermal dust emission with temperatures of $T_d \sim 35\text{--}40\text{ K}$ (Chapman et al. 2005; Kovács et al. 2006; Pope et al. 2006; Coppin et al. 2008), peaking in the FIR at rest frame $\lambda \sim 100 \mu\text{m}$. Due to the steep rise with frequency of the SED on the Rayleigh–Jeans tail ($S_{\nu} \propto \nu^{-3-4}$; Dunne et al. 2000; Dunne & Eales 2001), the FIR peak is increasingly redshifted into the sub-mm/mm observing bands with increasing distance, resulting in a strong negative k -correction that roughly cancels the effects of cosmological dimming with redshift for observations at $\lambda \gtrsim 500 \mu\text{m}$. This makes SMGs of a given bolometric luminosity equally detectable between $1 < z < 10$ at 1.1 mm, assuming that a sufficient amount of dust can be built up in these systems at such early epochs. Extragalactic surveys at sub-mm/mm wavelengths have long taken advantage of this unique property and have detected hundreds of starburst galaxies in the early Universe, many of which go undetected in even the deepest optical surveys due to extreme dust obscuration (e.g. Younger et al. 2007, 2009). These include deep, large-area surveys towards “blank fields”, i.e. regions devoid of known galaxy overdensities (e.g. Borys et al. 2003; Greve et al. 2004, 2008; Laurent et al. 2005; Coppin et al. 2006; Bertoldi et al. 2007; Perera et al. 2008; Scott et al. 2008; Weiß et al. 2009; Austermann et al. 2010) and towards “biased” fields, such as the environments of high-redshift radio galaxies (e.g. Stevens et al. 2003) and other tracers of high-redshift proto-clusters (e.g. Tamura et al. 2009). Several small-area surveys towards known clusters at moderate redshifts ($z \lesssim 0.5$) have also been carried out in order to detect intrinsically less luminous background SMGs via amplification through gravitational lensing (e.g. Smail et al. 1998, 2002; Chapman et al. 2002; Cowie, Barger & Kneib 2002; Knudsen et al. 2006; Knudsen, van der Werf & Kneib 2008).

The number counts at sub-mm/mm wavelengths provide potentially strong constraints on models of galaxy evolution (e.g. Granato et al. 2004; Baugh et al. 2005; Silva et al. 2005; Rowan-Robinson 2009). The hundreds of SMGs detected over the past 12 yr in large-area surveys taken with the Submillimeter Common-User Bolometer Array (SCUBA, 850 μm ; Holland et al. 1999) on the 15-m James Clerk Maxwell Telescope (JCMT), MAMBO (1.2 mm; Kreysa et al. 1998) on the Institut de Radio Astronomie Millimétrique (IRAM) 30-m telescope, Bolocam (1.1 mm; Glenn

et al. 1998; Haig et al. 2004) on the 10-m Caltech Submillimeter Observatory (CSO) and AzTEC (1.1 mm; Wilson et al. 2008) on the JCMT and the 10-m Atacama Submillimeter Telescope Experiment (ASTE, Ezawa et al. 2004, 2008) have put tight constraints on the number counts at $S_{850 \mu\text{m}} \geq 2\text{ mJy}$ ($S_{1.1 \text{ mm}} \geq 1\text{ mJy}$). Regardless of the precise details of various galaxy evolution models, the observed number counts at sub-mm/mm wavelengths require strong luminosity evolution of IR-bright galaxies (e.g. Scott et al. 2002; Greve et al. 2005; Coppin et al. 2006; Scott, Dunlop & Serjeant 2006; Austermann et al. 2010). The low angular resolution (full width at half-maximum, $\text{FWHM} \geq 10''$) of these observations makes them confusion-limited at $\sim 1\text{ mJy}$ (at both 850 μm and 1.1 mm on the 15-m JCMT), limiting source detections to only the most luminous ($L_{\text{FIR}} \gtrsim 3 \times 10^{12} L_{\odot}$) systems with extreme star formation rates of $\text{SFR} \gtrsim 500 M_{\odot}/\text{yr}$. Only surveys towards lensing clusters have picked out individual sources with lower luminosities, and these currently provide only weak constraints on the sub-mm number counts at $S_{850 \mu\text{m}} \lesssim 2\text{ mJy}$. While detecting a large number of intrinsically fainter systems and establishing the link between SMGs and other high-redshift galaxy populations must await the improved resolution of the 50-m Large Millimeter Telescope (LMT) and the Atacama Large Millimeter/Submillimeter Array (ALMA), information about their number counts at $S_{1.1 \text{ mm}} < 1\text{ mJy}$ can be gleaned from wide-area, confusion-limited surveys with existing facilities, and these can aid in discriminating between various galaxy evolution models.

In this paper, we present a 270 arcmin^2 1.1 mm survey of the Great Observatories Origins Deep Survey-South (GOODS-S) field taken with the AzTEC camera on the ASTE. This is the deepest survey at mm wavelengths ever carried out, achieving a root mean square (rms) noise level of $1\sigma = 0.48\text{--}0.73\text{ mJy}$. The GOODS-S field represents one of the most widely observed regions of sky, with deep multi-wavelength data from a number of ground-based and space-based facilities. This includes X-ray data from *Chandra* (Luo et al. 2008), optical to near-IR photometry from the *Hubble Space Telescope* (*HST*; Giavalisco et al. 2004), near-IR imaging with ISAAC on the Very Large Telescope (Retzlaff et al. 2010), *Spitzer* IRAC (Chary et al. in preparation) and MIPS (Dickinson et al. in preparation) imaging in the mid-IR, sub-mm imaging at 250, 350 and 500 μm with the Balloon-borne Large Aperture Submillimeter Telescope (BLAST; Devlin et al. 2009) and at 870 μm with the LABOCA camera on the Atacama Pathfinder EXperiment (APEX; Weiß et al. 2009) and 1.4 GHz interferometric imaging with the Very Large Array (VLA; Kellermann et al. 2008; Miller et al. 2008). Dedicated spectroscopic follow-up of optical sources in this field has also been underway (Vanzella et al. 2005, 2006, 2008; Popesso et al. 2009). Additionally, planned *Herschel* observations of the GOODS-S field will provide ultra-deep FIR/sub-mm imaging at 100, 250, 350 and 450 μm . This suite of

multi-wavelength data is essential for the identification of counterparts to the SMGs and for the characterization of the properties of these galaxies.

This paper is organized as follows. In Section 2, we describe the observations of the GOODS-S field carried out with AzTEC on the ASTE. In Section 3, we summarize the data reduction methods. We present the 1.1 mm map and source catalogue in Section 4 and describe simulations carried out to characterize the number of false detections, survey completeness and degree of source blending in the map in Section 5. We derive the 1.1 mm number counts from this survey in Section 6 and compare them with the number counts determined from SCUBA lensing cluster surveys and existing blank-field surveys at 1.1 mm wavelengths. We present a comparison between the AzTEC/GOODS-S data and the 870 μ m data from LABOCA in Section 7. We discuss the contribution to the cosmic infrared background (CIRB) at 1.1 mm from the radio and mid-IR galaxy populations in Section 8 and close with a summary of our results in Section 9. Several upcoming papers involving this data set are underway, including the identification of radio and mid-IR counterparts to the SMGs (Scott et al. in preparation), a study of the 1.1 mm properties of *BzK*-selected galaxies (Welch et al., in preparation), a comparison between the AzTEC/ASTE GOODS-S data and the sub-mm maps from BLAST (Aguilar et al., in preparation), and a study of the X-ray properties of SMGs (Johnson et al., in preparation).

2 OBSERVATIONS

We imaged a 26×20 arcmin² field centred at right ascension and declination (RA, Dec) = (03^h32^m30^s, $-27^\circ 48' 20''$) at 1.1 mm using AzTEC on the ASTE. The central 19×14 arcmin² region, where the coverage is uniform, encompasses the entire GOODS-S region mapped by the *Spitzer Space Telescope*. The observations were carried out using the N-COSMOS3 network observation system (Kamazaki et al. 2005) from July 15 to August 6 during the 2007 Chilean winter under excellent observing conditions, with $\tau_{220} = 0.05$ on average, and $\tau_{220} < 0.06$ 70 per cent of the time (zenith opacity at 220 GHz reported by the ASTE tau monitor). A total of 52 h of observing time excluding pointing and calibration overheads were devoted to this survey. During the 2007 season, 107 (out of 144) of the AzTEC bolometers were stable with high sensitivity. The point spread function (PSF) of each detector was measured via beam maps on Uranus, Neptune and 3C279 as described in Wilson et al. (2008) and has an FWHM of 30 ± 1 arcsec and 31 ± 2 arcsec in azimuth and elevation, respectively (the theoretical beam size is 27 arcsec FWHM). The full array subtends a circular field of view with a diameter of 8 arcmin.

2.1 Scan strategy

We used a continuous scanning strategy which traces a modified Lissajous pattern on the sky:

$$\begin{aligned}\Delta RA &= 5.5 \cdot \sin(a \cdot t + 0.25) + 2.0 \cdot \sin(a \cdot t / 30) \\ \Delta Dec &= 7.5 \cdot \sin(b \cdot t) + 2.0 \cdot \sin(b \cdot t / 30),\end{aligned}\quad (1)$$

where $a/b = 8/9$ and ΔRA and ΔDec are physical coordinates relative to the field centre. The actual values of a and b were scaled to limit the peak scanning velocity to 300 arcsec s^{-1} , and a rotational angle 20° west of north was used in order to align our map with that of the *Spitzer* IRAC/MIPS coverage of GOODS-S. A single observation took 42 min to complete, and we obtained a total of 74 observations of the GOODS-S field.

The benefit of Lissajous scanning, in addition to attaining excellent cross-linking and uniform coverage in the map, is that we avoid large telescope accelerations that can induce systematics in the detector signals as well as compromise the pointing accuracy of the telescope. Such effects are often seen in images taken in raster-scan mode, where $1/3 - 1/2$ of the data taken during times when the telescope reverses direction must be discarded (e.g. Perera et al. 2008; Scott et al. 2008). Lissajous scanning, on the other hand, results in nearly 100 per cent observing efficiency.

2.2 Pointing corrections

We make small corrections to the telescope pointing model by routinely observing the bright point source J0455–462 ($S_{1.1\text{ mm}} \sim 1.5$ Jy, variable) every 2 h before and after each block of GOODS-S observations. We measure pointing offsets by fitting the 4×4 arcmin² maps of J0455–462 to two-dimensional Gaussians, and we linearly interpolate these offset corrections temporally and apply them to the GOODS-S data. The random pointing error in the final GOODS-S map is $\lesssim 1$ arcsec (see Section 4.2).

2.3 Flux calibration

The flux conversion factor (FCF) used to convert the raw detector signals to flux density units was determined by beam map observations on Uranus, Neptune and 3C279, taken 1–2 times per night as described in Wilson et al. (2008). The flux densities of Uranus and Neptune at 1.1 mm were calculated from their frequency-dependent brightness temperatures reported in Griffin & Orton (1993) and ranged from 43 to 52 Jy and 18 to 20 Jy, respectively, during the AzTEC/ASTE 2007 observing season. The flux density of 3C279, which is highly variable, is adopted from the Submillimeter Array flux density archive¹ and ranged from 7.0 to 9.4 Jy at 1.1 mm during this time. We remove the responsivity factor from the detector signals and correct for extinction by modelling both as a linear function of the demodulated DC-level (see Wilson et al. 2008).

The measured FCF varied significantly from night to night, resulting in a 1σ scatter of 17 per cent over the entire observing run. We have identified the source of this scatter as the changing focal point of the telescope with environment temperature: the FCF decreases as the measured FWHM of the beam increases. Since real time corrections to the subreflector position were not possible, we use the same-night measurement of the FCF to calibrate each observation. To estimate the total calibration uncertainty, we determine the standard deviation in the measured flux densities from the 68 pointing observations of J0455–462, which is 8 per cent. Since this source is known to be variable, this gives a conservative upper limit to our calibration uncertainty. Combining this in quadrature with the 5 per cent absolute uncertainty on the flux densities of Uranus and Neptune (Griffin & Orton 1993) gives a total upper limit to the calibration error of 9 per cent.

3 DATA REDUCTION

We reduce the 1.1 mm data in a manner that is nearly identical to that described in detail in Scott et al. (2008). We summarize the steps here and note the differences. The raw time-stream data, which consist of all bolometer signals and pointing data stored as a function of time, are first scanned for “spikes” (defined as any $> 3\sigma$

¹<http://sma1.sma.hawaii.edu/callist/callist.html>

jump between sequential detector samples), which are caused by instrumental glitches or cosmic-ray strikes. These data and nearby samples, which amount to <0.1 per cent of the total time-stream data, are flagged and discarded from the data set. We group the remaining samples into 10-s intervals and then “clean” each 10-s group using a principal component analysis (PCA) to identify and remove the common-mode atmospheric signal (Laurent et al. 2005; Scott et al. 2008).

For AzTEC maps taken in raster-scan mode (e.g. Perera et al. 2008; Scott et al. 2008), data samples from the same individual scan (a single pass of the telescope across the sky) were grouped together for PCA cleaning. Since there is no such natural division for continuous Lissajous-scan maps, we chose a 10-s interval grouping. Perera et al. (2008) showed via a statistical correlation analysis that this PCA cleaning technique using time intervals ranging from 5 to 15 s provides a good balance between using a sufficient number of samples to determine the bolometer–bolometer correlations and being on a short enough time-scale so that the slower electronics-related low-frequency drifts can be effectively removed. We have verified that cleaning on 5–20-s intervals gives consistent results for AzTEC/GOODS-S.

After cleaning the time-stream data, the bolometer signals are calibrated and binned into 3×3 arcsec² pixels to make a map for each separate observation. These 74 maps are then co-added and Wiener filtered to suppress correlated large-scale structure from residual atmosphere and pixel-to-pixel variations on scales smaller than the beam in order to optimize the map for point-source detection (Scott et al. 2008). In addition to this filtered map, we track the effects

of PCA cleaning and filtering on a model point source (referred to hereafter as the point-source kernel). The smoothing slightly broadens the FWHM of the beam: fitting a two-dimensional Gaussian to the point-source kernel results in an FWHM of $34''.6$ and $34''.3$ in RA and Dec, respectively. We also generate 100 noise maps, each a realization of the signal-free noise in the GOODS-S map, by “jackknifing” the time-stream signals in the same manner described in Scott et al. (2008). The point-source kernel and noise realizations are used later in the analysis for simulation purposes (Section 5) and number counts determination (Section 6).

4 1.1 MM MAP AND SOURCE CATALOGUE

4.1 AzTEC 1.1 mm map

The full 512 arcmin^2 1.1 mm map of GOODS-S taken with AzTEC on the ASTE is shown in Fig. 1. Since we must consider only a region with uniform noise for most of the analysis presented in this paper, we define a 270 arcmin^2 “uniform coverage region”, where the coverage is greater than or equal to 50 per cent of the maximum coverage in the map. The 1σ rms noise in this region ranges from 0.48 to $0.73 \text{ mJy beam}^{-1}$, making this the deepest contiguous region ever mapped at 1.1 mm. The scanning strategy that we used to map this field results in two separate patches where the coverage is deepest ($0.48 \text{ mJy beam}^{-1}$), whereas the centre of the map is slightly shallower ($0.56 \text{ mJy beam}^{-1}$, see rms contours in Fig. 1). The noise determined from the jackknifed noise realizations is Gaussian distributed with $1\sigma = 0.57 \text{ mJy beam}^{-1}$.

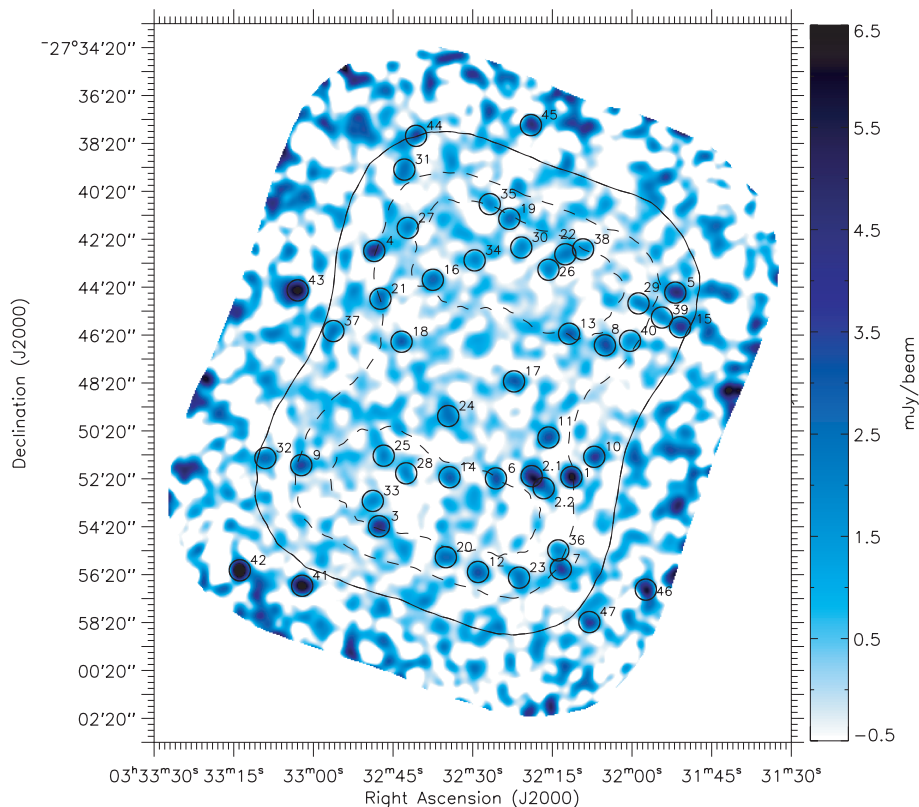


Figure 1. The AzTEC 1.1 mm map of the GOODS-S field. The solid contour shows the boundary of the 270 arcmin^2 50 per cent uniform coverage region to which most of the analysis of this field is restricted. The dashed contours (innermost to outermost) indicate noise rms levels of 0.51 and $0.57 \text{ mJy beam}^{-1}$. The circles located within the solid contour (diameter $= 2 \times \text{FWHM of AzTEC on ASTE} = 60''$) indicate the positions of $\geq 3.5\sigma$ source candidates, labelled in order of decreasing S/N of the detections. The circles located outside of the solid contour indicate robust $\geq 4.5\sigma$ sources within the noisier regions of the map. See the on-line journal for a colour version of this Figure.

4.2 Astrometry

To check for a residual astrometric offset after applying the pointing model (see Section 2.2), we stack the 1.1 mm GOODS-S map at the positions of radio sources in this field, whose positions are known to much better than 1 arcsec. From the VLA 1.4 GHz survey of the Extended Chandra Deep Field-South (ECDF-S; Miller et al. 2008), which reaches an rms noise level of 7–10 μ Jy, we extract a $\geq 4\sigma$ radio source catalogue² and stack the AzTEC map at the positions of the 219 radio sources that lie within the 50 per cent uniform coverage region. The resulting stacked AzTEC map shows a clear peak with a signal-to-noise ratio $S/N = 8$, and its centroid is offset by $-6''.3 \pm 2''.1$ in RA and $-3''.9 \pm 2''.1$ in Dec. We have verified this result by stacking at the locations of 1185 MIPS 24 μ m sources detected with $S/N > 10$ in the *Spitzer* GOODS-S survey (Dickinson et al., in preparation), which gives a $S/N = 11$ detection with a centroid offset of $-6''.3 \pm 1''.6$ in RA and $-0''.7 \pm 1''.6$ in Dec, consistent with the radio-stacking results. We thus apply an astrometric correction of $(\Delta RA, \Delta Dec) = (6''.3, 0''.7)$ to the AzTEC map and 1.1 mm source positions (we favour the offsets measured from the 24 μ m stacking due to the higher S/N of the peak detection).

This systematic offset represents the *average* pointing offset between J0455–462 and the GOODS-S field over 74 observations. The scatter in this offset, or the pointing jitter, will manifest itself as a broadening of the source in excess of the point-source kernel. We estimate the random pointing error in the AzTEC map from the brightest AzTEC source in this field, AzTEC/GS1. We adopt a simple model which consists of the convolution of the ideal point-source kernel with a two-dimensional Gaussian with standard deviation $(\sigma_{RA}, \sigma_{Dec})$, where σ_{RA} and σ_{Dec} are the 1σ random pointing errors in RA and Dec (see Scott et al. 2008, for a full description of this measurement). The high S/N of AzTEC/GS1 (11.6 σ , see Table 1) allows a clean measurement and provides a strong constraint on the random pointing error, since the signal from this single source can only be broadened with respect to the point-source kernel due to the scatter in the pointing model. The maximum likelihood estimate for the random pointing error from AzTEC/GS1 is $(\sigma_{RA}, \sigma_{Dec}) = (0''.5, 0''.1)$; however, the distribution of $(\sigma_{RA}, \sigma_{Dec})$ is very flat out to 1 arcsec, then falls off steadily. From this, we conclude that the random pointing error in the AzTEC/GOODS-S map is $\lesssim 1$ arcsec.

While the stacked 1.1 mm signal for the radio and MIPS 24 μ m populations is valuable for determining the astrometric offset, we find that the stacked signal cannot be used for determining the pointing jitter as it is significantly broader than the point-source kernel, with $1\sigma \approx 6.8$ arcsec and 15 arcsec for the stack on the radio and 24 μ m sources, respectively. This additional broadening is likely caused by other effects such as confused/blended sources in the 1.1 mm image or clustering of the radio and 24 μ m sources.

4.3 Source catalogue

We identify point sources in the 1.1 mm S/N map by searching for local maxima within 15 arcsec of pixels with $S/N \geq 3.5$ (see Scott et al. 2008, for a more detailed description). The 40 source candidates located within the uniform coverage region that meet

this criterion are listed in Table 1 in order of decreasing S/N of the detection. We note that the number of selected source candidates is independent of the window size for grouping high S/N pixels, for values ranging from 3 arcsec to 45 arcsec. Table 1 includes both the 1.1 mm flux densities and 1σ errors measured from the map, as well as the bias-corrected flux densities estimated using a semi-Bayesian technique (Section 6.2; Coppin et al. 2005, 2006; Austermann et al. 2009, 2010). This flux-bias correction accounts for the fact that the measured flux densities of mm-selected galaxies, which are generally detected at low S/N , are preferentially “boosted” due to the steep luminosity distribution of the population (e.g. Hogg & Turner 1998).

We find an additional seven robust source candidates that are located outside of the uniform coverage region, but are detected with high S/N (≥ 4.5). These sources are listed in Table 1 for the benefit of future studies; however, we do not address the general properties of these sources located outside of the uniform coverage region (e.g. number of false detections, completeness) nor include them in our number counts estimation in Section 6.

Several of the source candidates appear extended in the 1.1 mm map, most notably, AzTEC/GS2 (labelled as 2.1 and 2.2 in Fig. 1). To separate the components of AzTEC/GS2, we fit the 1.1 mm map in the neighbourhood of this source to a two-component model, where each component is a scaled version of the point-source kernel. The best-fit positions and flux densities are listed in Table 1. Given the comparatively modest S/N of the other AzTEC sources which are potentially extended, we cannot fit a model to these sources in order to separate them into multiple components.

5 CHARACTERISATION OF A DEEPLY CONFUSED MAP

Source confusion arising from a significant underlying population of faint sources that are individually undetected in a given survey is a function of source density and resolution (e.g. Hogg 2001, and references therein) and is important to consider for our AzTEC/GOODS-S survey. Using the standard rule of thumb (one source per 30 beams)³ and the blank-field 1.1 mm source counts from the AzTEC/SHADES survey (Austermann et al. 2010), the confusion limit given the 30 arcsec FWHM ASTE beam is 2.0 mJy. Since our AzTEC/GOODS-S survey is 3–4 times deeper than the formal confusion limit, we use specialized simulations to characterize the effects of source confusion on the properties of our GOODS-S map, including the number of false detections, the survey completeness, the positional uncertainty distribution and the degree of source blending.

For these purposes, we generate sky realizations of the GOODS-S field using the noise maps and simulated point sources, each modelled as the point-source kernel scaled by the source flux density. These simulated galaxies are injected at random positions drawn from a uniform distribution into the noise maps, where their flux density distribution is described by the best-fit Schechter function to the AzTEC/GOODS-S number counts (Section 6) truncated at $S_{1.1\text{ mm}} = 0.1$ mJy. These maps, referred to hereafter as “fully simulated maps”, provide a realistic model of the mm-galaxy population

²The radio source catalogue is produced using the SAD program in the Astronomical Image Processing System (AIPS; <http://www.aips.nrao.edu/>), which uses 2-dimensional Gaussian modelling to identify sources and produce photometry results.

³By deriving a general formalisation of source confusion statistics, Takeuchi & Ishii (2004) demonstrate the basis for the “one source per 30 beams” rule of thumb for estimating the confusion limit. Takeuchi & Ishii (2004) also show that the actual confusion limit is strongly dependant on the steepness of the source counts, rather than simply the source density.

Table 1. AzTEC/GOODS-S source candidates. The columns give: 1) AzTEC source identification; 2) source name; 3) S/N of the detection in the AzTEC map; 4) measured 1.1 mm flux density and error; and 5) deboosted 1.1 mm flux density and 68 per cent confidence interval (Section 6.2). The first 41 sources (AzTEC/GS1–AzTEC/GS40) are detected with $S/N \geq 3.5$ within the uniform coverage region of the AzTEC/GOODS-S map, where two are expected to be false positives. The seven additional $S/N \geq 4.5$ sources (AzTEC/GS41–AzTEC/GS47) are located outside the uniform coverage region.

AzTEC ID	Source name	S/N	$S_{1.1 \text{ mm}}$ (measured) (mJy)	$S_{1.1 \text{ mm}}$ (deboosted) (mJy)
AzTEC/GS1*	AzTEC_J033211.46–275216.0	11.6	6.6 ± 0.6	$6.3^{+0.5}_{-0.6}$
AzTEC/GS2	AzTEC_J033218.48–275221.8	11.4	6.0 ± 0.5	$5.7^{+0.5}_{-0.6}$
GS2.1*	AzTEC_J033218.99–275213.8		6.6 ± 0.5	$6.3^{+0.5}_{-0.5}$
GS2.2*	AzTEC_J033216.70–275244.0		4.0 ± 0.5	$3.7^{+0.5}_{-0.5}$
AzTEC/GS3*	AzTEC_J033247.86–275419.3	9.4	4.8 ± 0.5	$4.5^{+0.5}_{-0.5}$
AzTEC/GS4*	AzTEC_J033248.75–274249.5	8.6	5.0 ± 0.6	$4.6^{+0.6}_{-0.6}$
AzTEC/GS5*	AzTEC_J033151.81–274433.9	7.8	4.8 ± 0.6	$4.4^{+0.6}_{-0.6}$
AzTEC/GS6*	AzTEC_J033225.73–275219.4	6.7	3.4 ± 0.5	$3.1^{+0.5}_{-0.5}$
AzTEC/GS7*	AzTEC_J033213.47–275606.7	6.7	3.9 ± 0.6	$3.5^{+0.6}_{-0.6}$
AzTEC/GS8*	AzTEC_J033205.12–274645.8	6.6	3.5 ± 0.5	$3.1^{+0.5}_{-0.5}$
AzTEC/GS9	AzTEC_J033302.56–275146.1	6.5	3.6 ± 0.6	$3.2^{+0.6}_{-0.5}$
AzTEC/GS10*	AzTEC_J033207.13–275125.3	6.3	3.9 ± 0.6	$3.5^{+0.6}_{-0.6}$
AzTEC/GS11	AzTEC_J033215.79–275036.8	6.2	3.4 ± 0.6	$3.1^{+0.6}_{-0.6}$
AzTEC/GS12*	AzTEC_J033229.13–275613.8	6.2	3.3 ± 0.5	$2.9^{+0.5}_{-0.5}$
AzTEC/GS13	AzTEC_J033211.91–274616.9	6.2	3.1 ± 0.5	$2.8^{+0.5}_{-0.5}$
AzTEC/GS14	AzTEC_J033234.52–275216.4	6.0	3.0 ± 0.5	$2.6^{+0.5}_{-0.5}$
AzTEC/GS15*	AzTEC_J033150.91–274600.4	6.0	4.0 ± 0.7	$3.5^{+0.7}_{-0.7}$
AzTEC/GS16	AzTEC_J033237.67–274401.8	5.7	2.8 ± 0.5	$2.5^{+0.5}_{-0.5}$
AzTEC/GS17	AzTEC_J033222.31–274816.4	5.6	3.1 ± 0.6	$2.7^{+0.5}_{-0.6}$
AzTEC/GS18*	AzTEC_J033243.58–274636.9	5.5	3.1 ± 0.6	$2.7^{+0.5}_{-0.6}$
AzTEC/GS19	AzTEC_J033223.21–274128.8	5.4	2.7 ± 0.5	$2.4^{+0.5}_{-0.5}$
AzTEC/GS20	AzTEC_J033235.22–275536.8	5.2	2.7 ± 0.5	$2.4^{+0.5}_{-0.5}$
AzTEC/GS21	AzTEC_J033247.60–274449.3	5.0	2.9 ± 0.6	$2.4^{+0.6}_{-0.6}$
AzTEC/GS22	AzTEC_J033212.60–274257.9	4.7	2.3 ± 0.5	$2.0^{+0.5}_{-0.5}$
AzTEC/GS23*	AzTEC_J033221.37–275628.1	4.7	2.5 ± 0.5	$2.1^{+0.6}_{-0.5}$
AzTEC/GS24	AzTEC_J033234.76–274943.1	4.6	2.5 ± 0.5	$2.1^{+0.6}_{-0.6}$
AzTEC/GS25*	AzTEC_J033246.96–275122.4	4.4	2.2 ± 0.5	$1.8^{+0.5}_{-0.5}$
AzTEC/GS26	AzTEC_J033215.79–274336.6	4.4	2.1 ± 0.5	$1.8^{+0.5}_{-0.5}$
AzTEC/GS27	AzTEC_J033242.42–274151.9	4.3	2.3 ± 0.5	$1.8^{+0.5}_{-0.5}$
AzTEC/GS28	AzTEC_J033242.71–275206.8	4.3	2.1 ± 0.5	$1.7^{+0.5}_{-0.5}$
AzTEC/GS29	AzTEC_J033158.77–274500.9	4.1	2.2 ± 0.5	$1.8^{+0.5}_{-0.6}$
AzTEC/GS30	AzTEC_J033220.94–274240.8	4.1	2.0 ± 0.5	$1.7^{+0.5}_{-0.5}$
AzTEC/GS31	AzTEC_J033243.06–273925.6	4.1	2.5 ± 0.6	$2.0^{+0.6}_{-0.7}$
AzTEC/GS32	AzTEC_J033309.35–275128.4	4.1	2.8 ± 0.7	$2.1^{+0.7}_{-0.7}$
AzTEC/GS33	AzTEC_J033249.03–275315.8	4.1	2.0 ± 0.5	$1.6^{+0.5}_{-0.5}$
AzTEC/GS34	AzTEC_J033229.77–274313.1	4.0	2.0 ± 0.5	$1.6^{+0.5}_{-0.5}$
AzTEC/GS35	AzTEC_J033226.90–274052.1	4.0	2.0 ± 0.5	$1.6^{+0.5}_{-0.5}$
AzTEC/GS36	AzTEC_J033213.94–275519.7	3.7	2.1 ± 0.6	$1.5^{+0.6}_{-0.5}$
AzTEC/GS37	AzTEC_J033256.48–274610.3	3.7	2.5 ± 0.7	$1.8^{+0.7}_{-0.7}$
AzTEC/GS38	AzTEC_J033209.26–274245.5	3.6	1.8 ± 0.5	$1.4^{+0.5}_{-0.5}$
AzTEC/GS39*	AzTEC_J033154.34–274536.3	3.5	2.1 ± 0.6	$1.5^{+0.6}_{-0.6}$
AzTEC/GS40	AzTEC_J033200.38–274634.6	3.5	1.9 ± 0.6	$1.4^{+0.6}_{-0.6}$
AzTEC/GS41*	AzTEC_J033302.39–275648.4	8.0	7.1 ± 0.9	$6.3^{+0.9}_{-0.9}$
AzTEC/GS42*	AzTEC_J033314.19–275609.6	7.8	9.2 ± 1.2	$7.9^{+1.1}_{-1.4}$
AzTEC/GS43*	AzTEC_J033303.24–274428.3	6.9	6.7 ± 1.0	$5.7^{+1.0}_{-1.0}$
AzTEC/GS44*	AzTEC_J033240.84–273801.1	5.0	3.7 ± 0.7	$3.0^{+0.8}_{-0.8}$
AzTEC/GS45*	AzTEC_J033219.12–273734.1	4.8	4.8 ± 1.0	$3.6^{+1.0}_{-1.1}$
AzTEC/GS46*	AzTEC_J033157.38–275658.0	4.7	5.9 ± 1.2	$4.0^{+1.3}_{-1.4}$
AzTEC/GS47*	AzTEC_J033208.06–275819.4	4.5	3.8 ± 0.8	$2.8^{+0.9}_{-0.8}$

*These source candidates have probable 870 μm counterparts detected with $S/N \geq 3.7$ in the LABOCA map (Section 7.1).

in the GOODS-S field as observed by *AzTEC* by properly including the effects from our data reduction methods. We use these fully simulated maps throughout this paper to investigate various properties of our map.

One might ask if the assumption of a spatially uniform population of sources in our simulations is justified. To date there are only weak constraints on the clustering properties of bright SMGs derived from incomplete and/or small-number samples, with significant disagreement among various surveys (e.g. Blain et al. 2004; Greve et al. 2004; Scott et al. 2006; Weiß et al. 2009). Furthermore, bright sources like those probed in these surveys have low number density and so will not bias the map characteristics probed in our simulations. Rather, it is the faint, high-density, unexplored population of SMGs that can bias our estimates. Unfortunately, until better resolution for sub-mm/mm surveys is available (for example *AzTEC* on the LMT), we have no sufficient observational evidence to guide us. While these faint SMGs may have similar clustering properties as other high-redshift galaxy populations (e.g. LBGs, BzKs, EROs), it is likely that the projected two-dimensional clustering strength of SMGs is very different from optically selected high-redshift galaxies, since sub-mm/mm surveys sample a much larger volume of space. In the absence of observational constraints on the clustering of faint SMGs, we choose to model the source population with a uniform spatial distribution.

5.1 False detections

Given the modest S/N of the source candidates, we expect some fraction of the *AzTEC* sources in GOODS-S to be spurious. One method commonly used to estimate the number of false positives is to run the source-finding algorithm on the *negative* of the signal map. However, the PCA cleaning used for atmosphere removal leaves the signal map (and the point-source kernel) with a mean of zero. Every real source that increases the number of positive pixels in the map will thus also increase the number of negative pixels, so the number of false detections measured from the negative of the real map will be overestimated. Instead, we estimate the number of false detections by identifying the number of “sources” extracted from the 100 pure noise realizations. The number of false detections expected as a function of limiting S/N ratio is shown in Fig. 2 (solid curve, diamonds). At $\geq 3.5\sigma$, we expect $\sim 2/40$ sources in our catalogue (5 per cent) to be spurious. Above $\geq 4.25\sigma$, the number of *AzTEC* sources expected to be false positives is consistent with zero.

This estimate, however, provides only an upper limit to the number of spurious detections. In the real map, the negative bias in the pixel flux distribution from the addition of sources decreases the number of high-significance positive noise peaks in the map. This effect was first demonstrated for the *AzTEC*/GOODS-N survey (Perera et al. 2008) and is even more pronounced for our confusion-limited GOODS-S map. To demonstrate this effect, we run the source-finding algorithm on 600 fully simulated maps: for each *detected* source, we search within a 17 arcsec radius to identify the corresponding *input* source and detected sources that cannot be traced back to an input source are false positives. We choose a 17 arcsec search radius to ensure recovery of an input source, since a source detected at $S/N \geq 3.5$ has a probability < 0.1 per cent of being detected > 17 arcsec from its true position (Iverson et al. 2007). However, given the high source density at faint flux densities, there is a non-negligible probability that an intrinsically faint source will be located within 17 arcsec of a *random* position in the map, and so this calculation will underestimate the number of false positives.

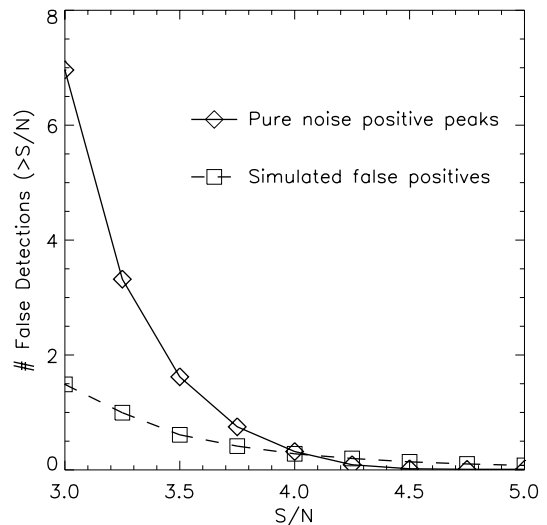


Figure 2. The expected number of false detections in the *AzTEC*/GOODS-S map as a function of limiting S/N. The solid curve and diamonds show the number of false detections estimated from the number of peaks detected in pure noise realizations with significance $\geq S/N$. The dashed curve and squares indicate the expected number of false positives determined from fully simulated maps. The pure noise realizations provide only a conservative upper limit to the number of false detections expected.

We account for the fraction of random associations as follows: (1) for each of the 600 fully simulated maps, we select 2000 random positions within the 50 per cent uniform coverage region, identify input sources located within 17 arcsec of these random positions, and from this determine the probability $P(> 17 \text{ arcsec} | S_i)$ that a source with intrinsic flux density S_i will be located within 17 arcsec of a random position in the map; (2) for each detected source in the simulated maps that can be traced back to an input source with intrinsic flux density S_i , we generate a random number, P_{test} , between 0 and 1 from a uniform distribution and consider the output–input source pair a false association (and hence the output source a false positive) if $P_{\text{test}} < P(> 17 \text{ arcsec} | S_i)$.

The number of false positives estimated from these fully simulated maps as a function of limiting S/N is shown in Fig. 2 (dashed curve, squares). From this estimate, we expect at most one of the $40 \geq 3.5\sigma$ sources in our catalogue to be spurious. At $\geq 3.0\sigma$, the number of false positives estimated from these fully simulated maps is significantly lower than that estimated from pure noise realizations (1–2 versus ~ 7), suggesting that we can comfortably extend our source catalogue to lower S/N detections. However, these simulations do not include the effects of source clustering. If the mm-galaxy population is strongly clustered on small angular scales, the strength of the negative bias in the pixel flux distribution would vary from region to region due to the variance in the source density, and thus the number of positive noise peaks (i.e. false positives) would be lower (higher) in the more (less) densely populated regions of the map. Since the clustering properties of the mm-galaxy population (including intrinsically faint sources) are not well known, we prefer to quote the values determined from the pure noise realizations as a conservative estimate of the number of false detections expected in our catalogue.

5.2 Completeness

The detection rate for a given source flux density is affected by both Gaussian random noise in the map and confusion noise from

the underlying bed of faint sources. To account for both effects, we estimate the survey completeness by measuring the recovery rate of point sources with known flux densities inserted into the real signal map, as described in Scott et al. (2008). For flux densities ranging from 0.1 to 8.0 mJy, we input 2000 sources per flux density *one at a time* into the GOODS-S map, each time randomly selecting the source position. Using the standard source-finding algorithm, an input source is considered recovered if it is detected in the map within 17 arcsec of its input position with a significance of $\geq 3.5\sigma$. We exclude samples where the simulated source was input or extracted less than 17 arcsec from a real $\geq 3.5\sigma$ source. The survey completeness is shown in Fig. 3 (data points with binomial error bars). The survey is 50 per cent complete at 2.1 mJy and 95 per cent complete at 3.5 mJy.

The sea of faint sources below the detection threshold adds confusion noise to the AzTEC/GOODS-S map. This additional noise reduces the map's sensitivity to individual sources and its survey completeness over a range of flux densities. An accurate estimate of the completeness is essential for correcting the observed number counts for this field. In the standard Bayesian method for extracting number counts from AzTEC maps (Austermann et al. 2009, 2010), survey completeness is estimated by injecting sources of known flux density one at a time into noise realizations and determining their recovery rate. While this method does not take confusion noise into account, this estimate was found to be consistent with that measured from the real signal map using the method described in the previous paragraph for several AzTEC surveys on the JCMT, where the angular resolution is better (FWHM = 18 arcsec on the JCMT versus 30 arcsec on the ASTE) and the 1σ depths of the surveys are ~ 1.0 mJy (e.g. Perera et al. 2008; Austermann et al. 2009, 2010), demonstrating that confusion noise was not significant for these observations. In contrast, we find that the completeness estimated from noise-only maps significantly overpredicts the survey com-

pleteness for our deeper, confusion-limited GOODS-S map (Fig. 3, dashed curve).

To verify that this difference arises from confusion noise, we next estimate the survey completeness from 10 000 fully simulated maps. For each $\geq 3.5\sigma$ source detected in these simulated maps, we identify the brightest input source within 17 arcsec of the output source position. We bin all detected sources by their input flux densities, and the completeness is calculated by the ratio of the number of recovered sources to the total number of input sources per flux bin. The completeness estimated from these fully simulated maps is shown as the solid curve in Fig. 3. This estimate agrees quite well with that from the single-input source simulations using the real GOODS-S map. The discrepancy between the two methods at $S_{1.1\text{mm}} \lesssim 1.5$ mJy likely arises from small imperfections in the assumptions we use to identify input–output pairs. For example, the single-input source simulations may slightly overestimate the completeness at low flux densities due to cases when the input source is inserted close to (but > 17 arcsec) a bright mm-galaxy in the real map. Still, despite the very different methods used for these two different completeness estimates, they differ by ≤ 4 per cent at all flux densities.

5.3 Positional uncertainty

The large beam combined with the low S/N of the detections results in a large positional error on the locations of sub-mm/mm-detected sources due to the effects of random and confusion noise in the map. We use the simulations described in 5.2, where a single source of known flux density is inserted into the GOODS-S map one at a time, to determine the distribution of input to output source distances as a function of detected S/N. The probability $P(>\theta; S/N)$ that a source will be detected outside of a radial distance θ of its true position is shown in Fig. 4 for three sample S/N bins. For comparison, the

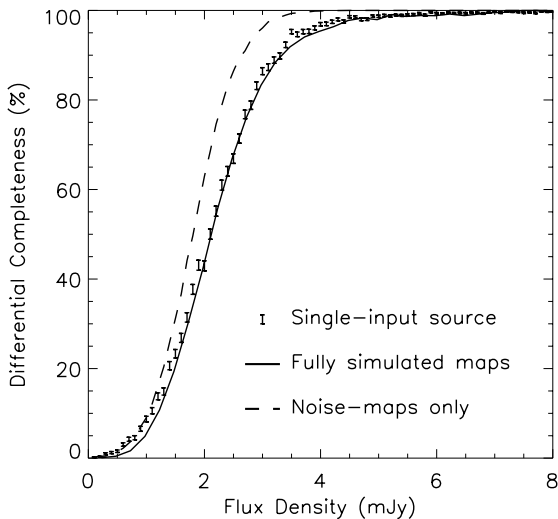


Figure 3. The survey completeness for $S/N \geq 3.5$ AzTEC sources in GOODS-S. The data-points and 68 per cent confidence binomial error bars show the completeness estimated by inserting sources of known flux density one at a time into the real signal map. The solid curve shows the completeness estimated by fully simulated maps. The dashed curve shows the completeness estimated by inserting sources of known flux density one at a time into pure noise realizations. This latter estimate does not account for the effects of confusion noise and hence results in an overestimate of the survey completeness for the range of flux densities shown here.

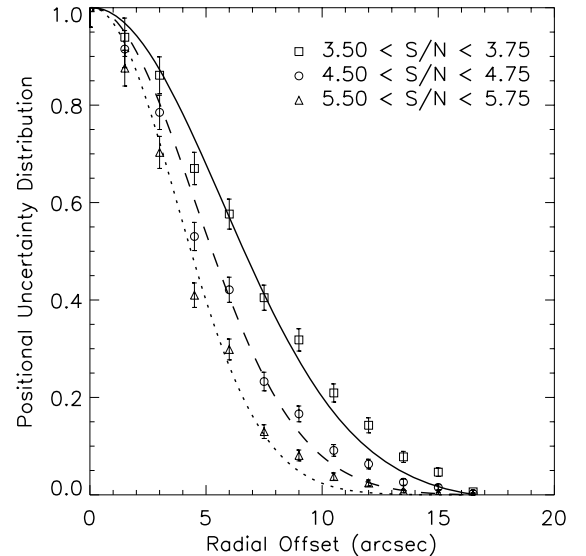


Figure 4. The positional uncertainty distribution for AzTEC/GOODS-S source candidates. The data-points and error bars show the probability $P(>\theta; S/N)$ that an AzTEC source detected with a given S/N ratio will be found outside a radial distance θ from its true location as determined from simulation (Section 5.3). The curves show the analytical expression derived in Iverson et al. (2007): solid, $3.5 < S/N < 3.75$; dashed, $4.5 < S/N < 4.75$; and dotted, $5.5 < S/N < 5.75$. The empirical and analytical distributions show broadly the same trend.

analytical solutions determined from Ivison et al. (2007) are shown as the solid ($3.5 < S/N < 3.75$), dashed ($4.5 < S/N < 4.75$) and dotted ($5.5 < S/N < 5.75$) curves, assuming that the FWHM of the AzTEC beam is 34 arcsec (i.e. the width of best-fitting Gaussian to the *filtered* point-source kernel). The analytical and empirical distributions show broadly the same trend.

5.4 Source blending

Given the depth and low angular resolution of the AzTEC/GOODS-S survey, some fraction of the $\geq 3.5\sigma$ sources in the map are expected to be the combined signal from two or more individual galaxies blended together. To estimate the fraction of the sources in Table 1 that are actually the blend of ≥ 2 individual galaxies, we take the $\geq 3.5\sigma$ sources detected from 600 fully simulated maps and trace each one back to all input sources located within 17 arcsec of the output source position. The fraction of detected sources that cannot be traced back to any input source is 0.8 per cent: these represent the ‘false positives’ estimated from fully simulated maps as described in Section 5.1. Sources that map back to only one input source are considered ‘single sources’, while those that can be traced back to two or more input sources are considered ‘blended sources’.

With this simple definition, we would expect nearly all sources to be blended given the large beam and the high source density of SMGs. However, a very faint source nearby a relatively bright source (for example, a 0.2 mJy source located 10 arcsec from a 3.0 mJy source) contributes a negligible amount to the summed signal. We want to avoid counting cases like these – where the brighter of the two sources completely dominates the total signal – as a blended source. As a more practical definition, we only consider a pair of nearby sources to be a blend if the contribution from each source to the summed signal is comparable. Using the input source flux densities and relative separations for all simulated sources within 17 arcsec of an output source, we model the beam-smoothed *noiseless* signal from the sum of these point sources and measure the peak flux density. If the fractional contribution to the summed flux density for an individual input source at the location of the peak is ≥ 70 per cent, we consider this a single source; otherwise, the detection is considered a blended source. With this definition, 25 per cent (10/40) of the AzTEC sources listed in Table 1 are expected to be ≥ 2 individual galaxies blended by the large beam. This fraction reduces to 15 per cent if we define a detection as a single source when it contributes ≥ 60 per cent to the summed noiseless flux density. These results represent lower limits to the fraction of blended sources, since this fraction would be even higher if the SMG population is significantly clustered. We note that by design, this fraction does not depend on the limiting S/N threshold (at least for $S/N \geq 3.5$).

5.5 Summary of confusion effects

We summarize our primary findings on the effects of confusion on the map properties.

(i) When the signal from confused sources becomes comparable to that from random noise fluctuations, the number of pure noise positive peaks in the map decreases. Thus, the number of false positives as a function of the limiting S/N is considerably lower than that estimated from peaks in noise-only realizations.

(ii) In terms of the survey completeness for a given catalogue, the confused signal acts as an additional noise component that reduces the map’s sensitivity to individual galaxies over a wide range in

intrinsic source flux density. It is important to include the effects of confusion when estimating the survey completeness for use in number counts calculations (Section 6.2).

(iii) A potentially large fraction of sources detected in a confusion-limited map are actually multiple galaxies blended by the large beam. Care must be taken to understand the possible biases this could cause in number counts estimations (Section 6.2).

6 NUMBER COUNTS

6.1 Fluctuation analysis

Due to the exceptional depth reached by this survey, the mm-emission from faint SMGs has a striking effect on the flux density distribution in the map. As discussed already, the method used to remove low-frequency modes leaves the mean of the map and the point-source kernel equal to zero, and every mm-source adds both positive and negative flux to the map. This is demonstrated in the left panel of Fig. 5, which shows the histogram of flux density values in the AzTEC/GOODS-S map (with Poisson error bars). The dashed curve shows the distribution of pixel fluxes averaged over 100 jackknifed noise realizations of this field and is very Gaussian. The flux distribution in the real map on the other hand is skewed by the presence of mm sources. While this effect makes the identification of individual galaxies challenging, we can use the distribution of flux values in the map to perform a fluctuation analysis, or more commonly referred to as a “ $P(d)$ ” analysis, in order to determine the number counts distribution for this field. As this technique is independent of the level of confusion in the map (at the expense of being somewhat model dependent), we can use it on the AzTEC/GOODS-S map to provide strong constraints on the SMG number counts at faint flux densities ($S_{1.1\text{ mm}} < 1\text{ mJy}$) well below the 1.1 mm confusion limit.

The fluctuation analysis is carried out as follows. Using a parameterized model of the number counts, we generate 100 simulated maps as described in Section 5 and compare the flux density distribution averaged over these simulated maps to that of the real GOODS-S map. For this single model, we calculate the comparison metric

$$-\ln(\mathcal{L}) = \sum_i m_i - d_i + d_i \cdot \ln(d_i/m_i), \quad (2)$$

where m_i represents the average number of pixels in the i th flux density bin from the model and d_i represents the corresponding quantity for the GOODS-S map. This process is repeated over a grid in parameter space to find the minimum of the above metric, which occurs at the best-fitting model. Minimizing this metric is equivalent to finding the maximum likelihood for the case that all histogram bins follow *independent* Poisson distributions. Note that we do not attempt to model effects of source clustering on the pixel flux distributions, since the clustering properties of the SMG population are not well known. This method is similar, in principle, to the parametric frequentist approach used by Perera et al. (2008) to determine number counts for the AzTEC/GOODS-N survey; however, here we consider the full flux density histogram in order to extract information about the faint source population.

For this analysis, we choose a Schechter function model given by

$$\frac{dN}{dS} = N_{3\text{ mJy}} \left(\frac{S}{3\text{ mJy}} \right)^{\alpha+1} \exp\left(\frac{-(S - 3\text{ mJy})}{S'} \right), \quad (3)$$

where $\frac{dN}{dS}$ is the differential number counts as a function of intrinsic 1.1 mm flux density S and $(S', N_{3\text{ mJy}}, \alpha)$ are the free parameters.

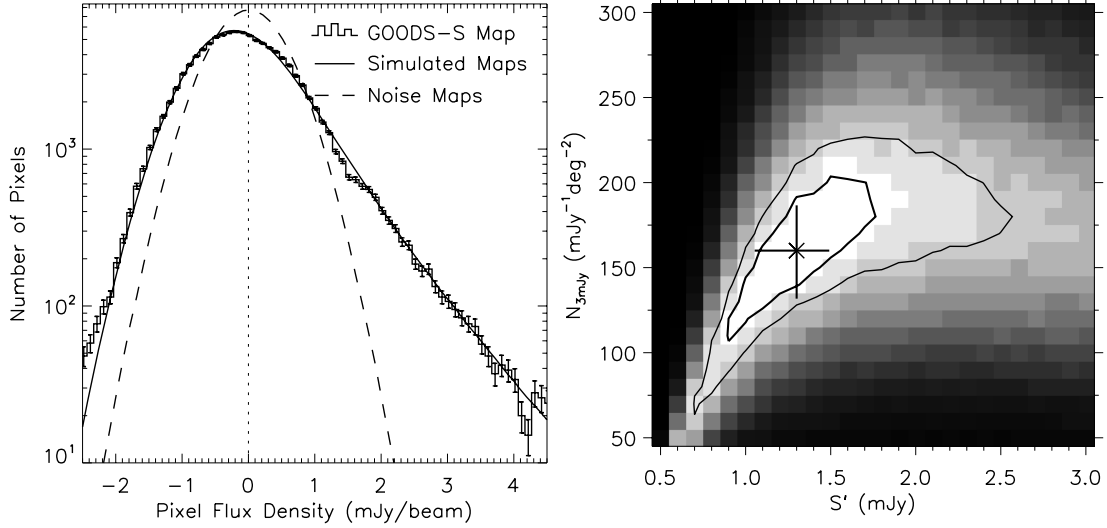


Figure 5. Left: The histogram of flux density values in the AzTEC/GOODS-S map. The dashed curve shows the distribution of flux values averaged over 100 noise realizations for this field, and is Gaussian distributed about zero. The solid curve shows the flux distribution averaged over fully simulated maps, populated according to the best-fit Schechter function model to the GOODS-S data. This demonstrates how a fluctuation analysis is used to determine the best model to the GOODS-S data. Right: The grey-scale shows the likelihood values of S' and $N_{3\text{mJy}}$ from the fluctuation analysis, and the cross indicates the best-fitting parameters to the data. The inner and outer contours indicate the 68.3 per cent and 95.5 per cent confidence regions, respectively. The error bars represent marginalized 68.3 per cent confidence intervals on each parameter.

While there are many forms of the Schechter function published in the literature, we prefer to fit to $N_{3\text{mJy}}$ because it reduces the degeneracies between the parameters, and it has been used in the number counts analyses of several previous AzTEC surveys (Perera et al. 2008; Austermann et al. 2009, 2010), making it straightforward to compare the results. We prefer a Schechter function model over that of a single power law because it allows for a natural steepening of the counts at high flux densities, which has been confirmed by large-area surveys at sub-mm/mm wavelengths (Coppin et al. 2006; Austermann et al. 2010) – though we are unlikely to see this steepening in the GOODS-S number counts given the small survey area. Since our data cannot provide strong constraints on a three-parameter fit, we fix $\alpha = -2$, a value that is consistent with estimates from previous surveys (e.g. Coppin et al. 2006; Perera et al. 2008; Austermann et al. 2009, 2010).

Since the Schechter function increases to infinity as S goes to zero, we must assume some minimum flux density cut-off, S_{\min} , for the population. A practical minimum flux limit is imposed by the data itself: at the flux density corresponding to where the number density of sources is ~ 1 per beam, adding fainter sources will not alter the flux density distribution in the map. Assuming the best-fitting model to the AzTEC/SHADES data (Austermann et al. 2010, which currently provides the tightest constraints on the blank-field 1.1 mm number counts), $S_{\min} \sim 0.1$ mJy. While it is not known whether the number density of sources for the SMG population turns over and starts to decrease somewhere below 1 mJy, the counts at the faint end of the 850 μm SCUBA galaxy population determined from lensing cluster surveys (e.g. Cowie et al. 2002; Smail et al. 2002; Knudsen et al. 2006, 2008) continue to rise out to $S_{850\mu\text{m}} \sim 0.2$ mJy, giving some reassurance that a 1.1 mm flux cut-off of $S_{\min} = 0.1$ mJy is reasonable. We use $S_{\min} = 0.1$ mJy in generating all simulated maps discussed in this paper; however, we have tested values ranging from $S_{\min} = 0.05$ to 0.3 mJy and have found that this does not affect the results from the fluctuation analysis.

Using 0.1 mJy bins for the flux histograms, we restrict the data model comparison to bins with ≥ 10 pixels on average (flux den-

ties between -2.8 and 5.5 mJy). The resulting best-fitting parameters are $(S', N_{3\text{mJy}}) = (1.30^{+0.19}_{-0.25} \text{ mJy}, 160^{+27}_{-28} \text{ mJy}^{-1} \text{ deg}^{-2})$. The flux distribution for this best-fitting model is shown as the solid curve in the left panel of Fig. 5. The likelihood values for the $S' - N_{3\text{mJy}}$ parameter space are shown in the right panel of Fig. 5, with the best-fitting parameters indicated by the cross. Due to the strong bin-to-bin correlations, it is not possible to determine the errors on the best-fitting parameters analytically. Instead, we determine the errors statistically through simulation by generating 600 fully simulated maps populated assuming the best-fitting Schechter function model to the real GOODS-S map (including Poisson deviations) and then performing the same fluctuation analysis on these simulated maps. The distribution of best-fitting parameters from these simulated maps is used to determine the 68.3 per cent and 95.5 per cent confidence intervals (contours in Fig. 5). The errors given for S' and $N_{3\text{mJy}}$ above (and shown as error bars in Fig. 5) represent the marginalized 68.3 per cent confidence intervals on each parameter.

The ability to recover the input number counts determined from the fully simulated maps verifies the reliability of this fluctuation analysis method. The best-fitting model to the actual GOODS-S data is listed in the first row of Table 2, and the differential number

Table 2. The best-fitting parameters for models to the AzTEC/GOODS-S number counts. The method used is listed in the first column: “ $P(d)$ ” = fluctuation analysis, and “Bayes” = Bayesian method. The errors on the best-fitting parameters represent marginalised 68.3 per cent confidence intervals. When an error is not listed, the parameter was fixed to the given value.

Method	Model	S' (mJy)	$N_{3\text{mJy}}$ (mJy $^{-1}$ deg $^{-2}$)	α
$P(d)$	equation (3)	$1.30^{+0.19}_{-0.25}$	160^{+27}_{-28}	-2.0
Bayes	equation (3)	$1.47^{+0.40}_{-0.25}$	131^{+30}_{-20}	-2.0
$P(d)$	equation (4)	—	90^{+20}_{-18}	$-3.70^{+0.18}_{-0.11}$
Bayes	equation (4)	—	107^{+30}_{-10}	$-3.35^{+0.25}_{-0.15}$

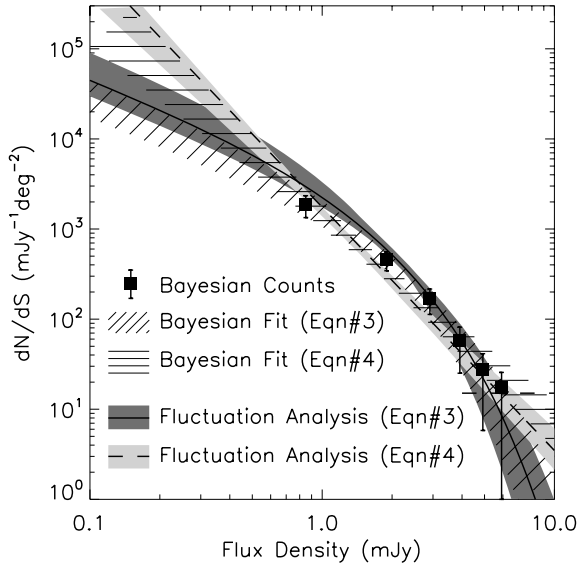


Figure 6. The differential number counts for the AzTEC/GOODS-S field using various methods and models. The solid curve and dark shaded region indicate the best-fitting model and 68.3 per cent confidence interval from the fluctuation analysis, assuming a Schechter function model. The dashed curve and light shaded region indicate the best-fitting model and 68.3 per cent confidence interval from a fluctuation analysis assuming a single power-law model. The squares show the Bayesian-extracted number counts and their 68.3 per cent confidence intervals determined in Section 6.2. The slanted (horizontal) hatching shows the 68.3 per cent confidence interval for a fit to the Bayesian-extracted counts, assuming a Schechter function (power-law) model for the population. The horizontal dashed line shows the survey limit, where the number counts will Poisson deviate to $0 \text{ mJy}^{-1} \text{ deg}^{-2}$ 31.7 per cent of the time given the area of the GOODS-S survey.

counts from this fluctuation analysis is shown in Fig. 6 as the thick solid curve (best-fitting) and dark shaded region (68.3 per cent confidence interval). These results are also shown in Figs 7 and 8 for a comparison with other surveys.

While potentially providing tight constraints on the source counts of the SMG population, this fluctuation analysis relies strongly on the accuracy of our noise realizations and the point-source kernel and is sensitive to the assumed number counts model. To demonstrate this, we carry out a fluctuation analysis on the GOODS-S data assuming instead a single power-law model for the number counts, which is given by

$$\frac{dN}{dS} = N_{3\text{mJy}} \left(\frac{S}{3\text{ mJy}} \right)^{\alpha+1}. \quad (4)$$

We chose this functional form to make both parameters easily comparable to those from the Schechter function model, since $N_{3\text{mJy}}$ gives the differential number counts at 3 mJy, and α represents the power-law dependence at low flux densities. The best-fitting parameters are listed in the third row of Table 2, and the best-fitting model and 68.3 per cent confidence interval are shown as the dashed curve and light shaded region in Figs 6–8. The best-fitting values of $N_{3\text{mJy}}$ for the two different models are only marginally consistent, and the best-fitting power-law index ($\alpha + 1 = -2.7$) is significantly steeper than that assumed for the Schechter function model ($\alpha + 1 = -1.0$). Comparing the two models in Fig. 6, they disagree significantly at $S \sim 2 \text{ mJy}$, where supposedly the tightest constraints can be placed on the number counts. One must thus exercise caution when interpreting results from this fluctuation analysis, as it inherently as-

sumes that the model provides a good representation of the source counts.

6.2 Bayesian estimation

We can minimize the dependence of the analysis on our underlying model of the number counts by utilizing a modified form of the semi-Bayesian method introduced by Coppin et al. (2005, 2006). This semi-Bayesian method is now widely used to extract the number counts from sub-mm/mm surveys because of its ability to handle various survey biases, and it has been extensively tested and validated using previous AzTEC data sets (Perera et al. 2008; Austermann et al. 2009, 2010). Since this method is described in detail in the aforementioned papers, we only briefly summarize the steps here.

The raw source counts from sub-mm/mm surveys suffer from three main biases: the “flux boosting” effect described in Section 4.3, contamination from false positives and incompleteness. To account for the first two effects, we generate posterior flux distributions (PFDs), $p(S_i | S_m, \sigma_m)$ (where S_i is the intrinsic flux density of the source, S_m is the measured flux density and σ_m is the error on the measured flux density), for each source candidate assuming some prior model for the SMG number counts. These PFDs are then randomly sampled (with replacement) to determine *intrinsic* flux densities for the sources, and these fluxes are binned to calculate differential and cumulative number counts. This process is repeated 20 000 times to adequately sample the number counts distribution. We also include sample variance by Poisson deviating the number of sources sampled in each of the 20 000 iterations. Since the PFD for each source candidate includes a non-negligible probability that the intrinsic flux density is $S_i < 0 \text{ mJy}$, this procedure inherently accounts for false positives.

To extract source counts from the AzTEC/GOODS-S map, we use the best-fitting Schechter function model determined from the fluctuation analysis as the prior distribution to generate PFDs for the source candidates. We sample all source candidates where the probability of the source having a flux density less than zero is $p(S_i < 0 | S_m, \sigma_m) \leq 0.05$, which corresponds to $S/N \gtrsim 2.8$ (actually depends on both S_m and σ_m) for a total of 54 source candidates. Austermann et al. (2010) tested various limiting thresholds for $p(S_i < 0 | S_m, \sigma_m)$ using several AzTEC data sets and found that any variations in the resulting number counts are much smaller than the formal 68.3 per cent uncertainties. They found that for accurate PFDs, the $p(S_i < 0 | S_m, \sigma_m)$ limiting threshold does not greatly affect the resulting number counts (provided source confusion is not an issue) and that using a higher limiting threshold supplies more information (due to increased survey completeness at faint flux densities) without introducing significant biases. We have verified that the PFDs for GOODS-S source candidates with $S/N \geq 2.8$ are accurate to better than 1 per cent at $S_i \geq 0.5 \text{ mJy}$ following the simulations described in Austermann et al. (2009, 2010), further justifying the use of a $p(S_i < 0 | S_m, \sigma_m) \leq 0.05$ limiting threshold in this analysis. We note that while Austermann et al. (2010) use a limiting threshold of $p(S_i < 0 | S_m, \sigma_m) \leq 0.2$, we choose a more conservative limit to avoid potential systematics arising from confusion effects, which were negligible for Austermann et al. (2010).

The raw number counts must also be corrected for incompleteness. In previous implementations of the semi-Bayesian method on AzTEC data sets (Perera et al. 2008; Austermann et al. 2009, 2010), survey completeness was estimated by the recovery rate of synthetic point sources with known intrinsic flux densities inserted into pure noise realizations one at a time. However, we have shown

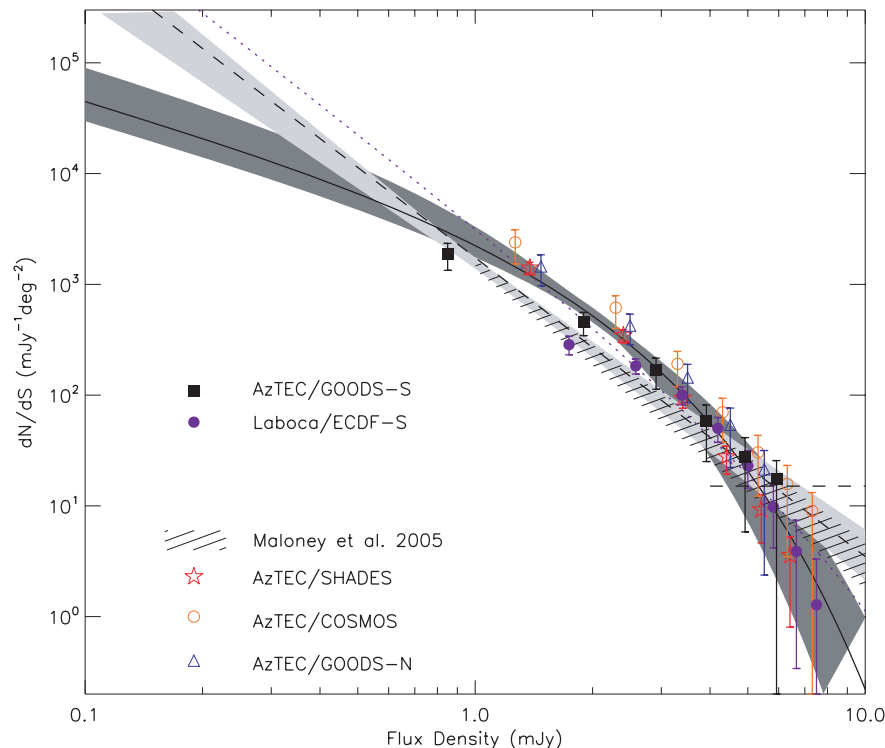


Figure 7. The differential number counts from the AzTEC/GOODS-S survey (*black squares*), compared with those determined from AzTEC surveys of other fields, including: GOODS-N (Perera et al. 2008), COSMOS (Austermann et al. 2009), and SHADES (Austermann et al. 2010). The error bars represent 68.3 per cent confidence intervals on Bayesian-extracted counts. The solid (dashed) curve and dark (light) shaded region indicate the best-fitting Schechter function (power-law) model and 68.3 per cent confidence region from a fluctuation analysis of GOODS-S (Section 6.1). The horizontal dashed line shows the survey limit, where the number counts will Poisson deviate to $0 \text{ mJy}^{-1} \text{ deg}^{-2}$ 31.7 per cent of the time given the area of the GOODS-S survey. The 95 per cent confidence interval from a fluctuation analysis of the 1.1 mm Bolocam Lockman Hole survey (Maloney et al. 2005) assuming a single power-law model is shown by the hatched region. The filled circles indicate the differential number counts determined from a Bayesian analysis of the $870 \mu\text{m}$ LABOCA survey of the ECDF-S, scaled to 1.1 mm assuming a flux ratio of $S_{870 \mu\text{m}}/S_{1.1 \text{ mm}} = 2.0$. The dotted curve shows the best-fitting Schechter function model to the LABOCA/ECDF-S data from a $P(d)$ analysis scaled to 1.1 mm (confidence region not available). See the on-line journal for a colour version of this Figure.

in Section 5.2 that this method overestimates the completeness for the AzTEC/GOODS-S field, since (unlike the AzTEC/JCMT surveys) confusion noise is significant for our data. Using a pure noise completeness estimate would consequently underestimate the number counts in this field. We instead estimate the survey completeness through fully simulated maps as described in Section 5.2, where the simulated maps are populated according to the number counts distribution given by the assumed *prior*. To be consistent with our catalogue selection from the real GOODS-S map, an input source in these simulations is considered to be recovered only if its corresponding output source passes the limiting threshold test of $p(S_i < 0 | S_m, \sigma_m) \leq 0.05$. Since we are using an “ideal” prior determined from the fluctuation analysis of this field, we are confident that this provides a good completeness estimate for the correction of the raw number counts.

The 1.1 mm differential number counts for the GOODS-S field determined from the Bayesian method are shown in Figs 6 and 7 (filled squares), and the cumulative number counts are shown in Fig. 8. The counts are also listed in Table 3. The number counts are calculated from the mean number of sources in each flux bin (with bin size = 1 mJy) over the 20 000 iterations, and the errors represent the 68.3 per cent confidence intervals calculated from the distribution in the counts across those iterations. For the differential number counts, the flux densities in Table 3 are the effective bin centres weighted by the assumed prior. The number counts from this method are highly correlated, since they are estimated by av-

eraging over many realizations of the number counts bootstrapped off the *same* source catalogue. The covariance matrices for the differential and cumulative number counts are listed in Tables 4 and 5, respectively.

The Bayesian-extracted number counts agree within the 1σ errors with the best-fitting Schechter function model from the fluctuation analysis. However, the number counts in the two lowest flux density bins (0.5–1.5 mJy and 1.5–2.5 mJy) are low compared to the best-fitting curve from the fluctuation analysis, while the number counts in the highest bin (5.5–6.5 mJy) are high. This may arise from a systematic bias in the Bayesian-extraction method due to source blending, since this technique does not account for the possibility that an individually detected source is the summed flux density of two or more galaxies. This would indeed result in the apparent bias seen here, as the number counts would be overestimated at high flux densities and underestimated at low flux densities. We have checked for this potential bias in the Bayesian method by running this analysis on source catalogues extracted from 600 fully simulated maps populated according to the best-fitting model from the fluctuation analysis. Since in this case we know the exact form of the source counts input into each map, we can search for such effects in the output number counts. We find that the output differential number counts for these simulated maps do indicate that this bias due to source blending is present in the data: for the 5.5–6.5 mJy flux density bin, the extracted number counts are higher than the input number counts for 65 per cent of the simulated maps, while for

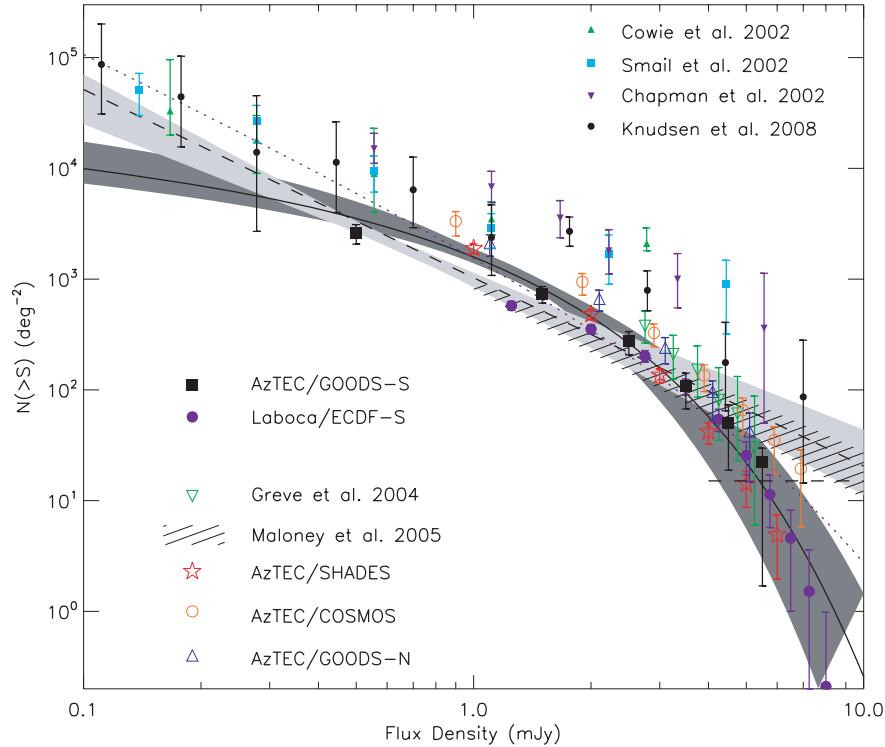


Figure 8. The cumulative number counts from the AzTEC/GOODS-S survey and AzTEC surveys of other fields. See the caption of Fig. 7 for a full description of the symbols. The 1.2 mm number counts from Greve et al. (2004) are shown as empty upside-down triangles with 95 per cent confidence error bars. For comparison, the number counts from SCUBA 850 μm lensing cluster surveys are shown with smaller symbols. All 850 μm counts have been scaled to 1.1 mm assuming a simple flux ratio scaling of $S_{850\mu\text{m}}/S_{1.1\text{mm}} = 1.8$. The number counts from Cowie et al. (2002), Smail et al. (2002), Chapman et al. (2002) and Greve et al. (2004) have not been corrected for flux boosting and therefore are likely biased to the right. See the on-line journal for a colour version of this Figure.

Table 3. The differential and cumulative number counts for the AzTEC/GOODS-S field, calculated using the Bayesian method described in Section 6.2. The errors indicate 68.3 per cent confidence intervals.

Flux Density (mJy)	dN/dS ($\text{mJy}^{-1} \text{deg}^{-2}$)	Flux Density (mJy)	$N(>S)$ (deg^{-2})
0.85	1892^{+453}_{-554}	0.50	2631^{+478}_{-562}
1.90	461^{+97}_{-116}	1.50	739^{+117}_{-132}
2.91	170^{+46}_{-57}	2.50	279^{+57}_{-72}
3.92	58^{+23}_{-33}	3.50	109^{+33}_{-41}
4.92	28^{+14}_{-22}	4.50	50^{+23}_{-31}
5.92	17^{+8}_{-17}	5.50	22^{+7}_{-21}

the 0.5–1.5 mJy bin, the extracted counts are lower than the input counts for 60 per cent of the simulated maps. However, at all flux densities, the extracted number counts agree with the input number counts within their 1σ (2σ) errors at least 86 per cent (96 per cent) of the time, so this bias is small compared to the formal Poisson errors.

To verify that the Bayesian-extracted number counts are insensitive to the assumed prior, we reran this procedure on the AzTEC/GOODS-S map using a prior distribution that is consistent with a fit to the number counts from the AzTEC/SHADES survey: $(S', N_{3\text{mJy}}, \alpha) = (1.11 \text{ mJy}, 153 \text{ mJy}^{-1} \text{deg}^{-2}, -2.0)$. We find that the extracted number counts using these two different priors agree within 4 per cent at flux densities $\geq 1.5 \text{ mJy}$. For the lowest flux bin at 0.5–1.5 mJy, the results agree within 19 per cent, i.e. well

Table 4. The covariance matrix for the differential number counts for the AzTEC/GOODS-S field. The units are in $\text{mJy}^{-2} \text{deg}^{-4}$.

Flux Density (mJy)	0.85	1.90	2.91	3.92	4.92	5.92
0.85	254800	37340	3559	82	-79	11
1.90		11420	2835	345	8	8
2.91			2623	848	73	5
3.92				808	353	27
4.92					366	127
5.92						230

Table 5. The covariance matrix for the cumulative number counts for the AzTEC/GOODS-S field. The units are in deg^{-4} .

Flux Density (mJy)	0.50	1.50	2.50	3.50	4.50	5.50
0.50	271300	16060	4350	1478	659	307
1.50		15580	4144	1470	650	300
2.50			4123	1486	667	293
3.50				1471	667	297
4.50					670	301
5.50						297

within the formal 1σ error. This demonstrates that the results from this technique are robust given a reasonable assumption for the prior number counts distribution. For this reason, we can fit these number counts to various models.

For a given model, we fit to each of the 20 000 bootstrap iterations separately as the flux density bins for a given iteration are uncorrelated (see Austermann et al. 2010), and we use the distribution of best-fitting parameters to determine the most likely values and their confidence intervals. The results of a fit to the GOODS-S differential number counts assuming a Schechter function model (equation 3) and a power-law model (equation 4) are given in rows 2 and 4 of Table 2 and are shown in Fig. 6 as the hatched regions, which indicate the 68.3 per cent confidence intervals. The fits to the Bayesian-extracted counts are in good agreement with the results from the fluctuation analysis for a given model. To compare the two models, we compute the χ^2 metric given by

$$\chi^2 = (\mathbf{d} - \mathbf{m})\mathbf{w}(\mathbf{d} - \mathbf{m})^T, \quad (5)$$

where \mathbf{d} is the row vector containing the differential number counts from the Bayesian method, \mathbf{w} is the corresponding weight matrix calculated from the inverse of the covariance matrix and \mathbf{m} is the model number counts evaluated at the same flux density bins as \mathbf{d} . For the best-fitting Schechter function to the Bayesian-extracted number counts, $\chi^2 = 0.84$. This model provides a better fit to the data than the single power-law model, for which $\chi^2 = 5.4$. Note that since the errors are not Gaussian-distributed, this metric is not expected to follow the χ^2 -distribution; we use this metric simply to compare the relative goodness of fit for these two models.

6.3 Comparison with results from SCUBA lensing cluster surveys

The AzTEC/GOODS-S survey currently provides the only constraints on the 1.1 mm number counts at $S_{1.1\text{ mm}} < 1\text{ mJy}$. For comparison, none of the existing blank-field SCUBA surveys constrains the 850 μm counts to comparatively faint flux densities ($S_{850\text{ }\mu\text{m}} \lesssim 2\text{ mJy}$): the deepest (e.g. Hughes et al. 1998; Eales et al. 2000) are too small in area to provide statistically significant samples of faint SMGs, while the large-area surveys (e.g. Borys et al. 2003; Coppin et al. 2006) do not reach sufficient depths to probe such faint sources. The most significant constraints on the $S_{850\text{ }\mu\text{m}} \lesssim 2\text{ mJy}$ number counts come from SCUBA lensing cluster surveys, where massive foreground clusters are used to probe faint background SMGs via gravitational lensing (e.g. Chapman et al. 2002; Cowie et al. 2002; Smail et al. 2002; Knudsen et al. 2006, 2008). This is a powerful technique for detecting intrinsically faint high-redshift SMGs and for estimating their source counts at low flux densities, as the foreground clusters magnify both the flux of the background sources (by factors of typically 2–3) and the area in the source plane, effectively decreasing the survey confusion-limit that hinders the sensitivity of blank-field observations. With new constraints on the $S_{1.1\text{ mm}} = 0.5\text{ mJy}$ number counts from the AzTEC/GOODS-S survey, the faint end of the number counts determined from lensing cluster surveys can be compared to the results from a blank field. Given the small sample size from lensing cluster surveys, this comparison is limited to the cumulative number counts.

The number counts from four separate SCUBA lensing cluster surveys are shown in Fig. 8. To compare these with our 1.1 mm results, we must account for the difference in observing wavelengths. We assume here that SCUBA and AzTEC are sampling the same underlying population of sub-mm/mm-bright galaxies and that the difference in the observed number counts can be described by a simple flux scaling: $R = S_{850\text{ }\mu\text{m}}/S_{1.1\text{ mm}}$. In principle, surveys at 1 mm may preferentially select sources at higher redshifts – or sources with colder dust temperatures – than surveys at 850 μm .

While there is some evidence that 1 mm surveys select on average higher redshift galaxies than 850 μm surveys (Eales et al. 2003; Younger et al. 2007, 2009; Greve et al. 2008), other studies suggest that there is no significant difference between the two populations (Greve et al. 2004; Ivison et al. 2005; Bertoldi et al. 2007).

A recent source-to-source comparison of SCUBA and AzTEC sources in the GOODS-N field (Chapin et al. 2009) reveals that while the redshift distribution of 1.1 mm sources in that field peaks at a higher redshift than that of 850 μm sources ($z = 2.7$ versus $z = 2.0$), the population is consistent with an average flux scaling with $R = 1.8$. Alternatively, the 850 μm number counts determined from the SCUBA/SHADES survey (Coppin et al. 2006) and the 1.1 mm number counts from the GOODS-N field (Perera et al. 2008) are consistent assuming a flux scaling of $R = 2.1 \pm 0.2$. These estimates are equivalent to the expected flux ratio of a $z = 2.5$ galaxy whose SED can be modelled as a single-component-modified black-body with $T_d = 30\text{ K}$ and emissivity index $\beta = 1.5$. As this model is consistent with the expected SEDs of local galaxies observed with the *Infrared Astronomical Satellite* and SCUBA (Dunne et al. 2000; Dunne & Eales 2001) as well as the measured SEDs of several SMGs (Chapman et al. 2005; Kovács et al. 2006; Pope et al. 2006; Coppin et al. 2008), we start by adopting a simple scaling factor of $R = 1.8$ for the purposes of comparing the number counts from the SCUBA lensing cluster surveys to the AzTEC/GOODS-S number counts.

The scaled number counts from the SCUBA lensing cluster surveys agree only marginally with the number counts from the AzTEC/GOODS-S survey and are systematically higher at all flux densities (see Fig. 8). In fact, the lensing cluster survey number counts are systematically higher than those from *all* AzTEC blank-field surveys. We first examine the possibility that the flux scaling factor of $R = 1.8$ derived from the GOODS-N field by Chapin et al. (2009) is not universal by deriving the value of R that minimizes the residuals between the SCUBA lensing cluster number counts and the best-fitting model to the 1.1 mm number counts from the AzTEC/GOODS-S field. The best-fitting values and 68.3 per cent confidence errors for R determined from each of the four lensing cluster data sets shown in Fig. 8 are given in Table 6. We find $R \geq 3.3$ for all lensing cluster surveys considered here, regardless of which model (Schechter function or power-law) we consider. Such high values of R are inconsistent with observed values for

Table 6. The 850 μm to 1.1 mm flux ratio estimated by scaling the number counts from SCUBA lensing cluster surveys to the best-fitting model to the AzTEC/GOODS-S number counts. The columns are: 1) reference for the lensing cluster survey (last row is the results for all four surveys combined); 2) best-fitting flux ratio $R = S_{850\text{ }\mu\text{m}}/S_{1.1\text{ mm}}$ and 68.3 per cent confidence interval assuming the Schechter function model for the GOODS-S number counts (equation 3); and 3) best-fitting flux ratio and 68.3 per cent confidence interval assuming the power-law model for the GOODS-S number counts (equation 4).

Reference	R (Schechter)	R (Power-law)
Cowie et al. (2002)	$5.6^{+0.9}_{-0.9}$	$3.5^{+0.4}_{-0.5}$
Smail et al. (2002)	$4.4^{+0.9}_{-0.8}$	$3.7^{+0.3}_{-0.4}$
Chapman et al. (2002)	$5.2^{+0.8}_{-0.7}$	$5.5^{+0.5}_{-0.5}$
Knudsen et al. (2008)	$3.3^{+0.4}_{-0.4}$	$3.3^{+0.4}_{-0.5}$
($S_{850\text{ }\mu\text{m}} \leq 2\text{ mJy}$)	$3.5^{+1.8}_{-1.4}$	$2.8^{+0.6}_{-0.7}$
All	$4.0^{+0.3}_{-0.3}$	$3.5^{+0.2}_{-0.3}$

individual bright SMGs (when both 850 μm and 1.1 mm measurements are available) as well as predicted values assuming SEDs and redshifts typical of this population. We note, however, that a high value of $R = 2.5 \pm 0.1$ has been estimated from a similar scaling of the 850 μm and 1.1 mm number counts from the SCUBA and AzTEC/SHADES surveys (Austermann et al. 2010). If 1.1 mm surveys are tracing a population of galaxies with either higher redshifts or colder dust temperatures than those detected at 850 μm , we would expect this flux ratio to be *lower*, since these two bands probe closer to the dust peak. The high value of R measured from scaling the SHADES number counts is likely due to systematics caused by differences in the number counts analyses or calibration, and we consider this an upper limit to the true flux ratio. Estimates of R from scaling the 850 μm lensing cluster number counts to the best-fitting model to the AzTEC/GOODS-S data set represent $\gtrsim 2\sigma$ deviations from this upper limit of $R = 2.5$ from the SHADES results, implying that even greater systematics exist in the lensing cluster number counts. Evidence for significant systematics in the derivation of the lensing cluster number counts is also given by the fact that the cumulative number counts at $S_{850\mu\text{m}} > 2\text{ mJy}$ from lensing cluster surveys are also higher than those from large-area blank-field SCUBA surveys (e.g. Borys et al. 2003; Webb et al. 2003; Coppin et al. 2006; Scott et al. 2006).

There are several issues in the number counts extraction from lensing cluster surveys which may systematically bias the number counts high. Errors in the detailed mass models of the lensing clusters are an obvious candidate for the observed bias; however, any errors in estimating magnification factors for source flux densities are compensated by equivalent errors in the source plane area (Knudsen et al. 2006, 2008). Contamination by foreground cluster members is another concern, but these have been identified and excluded in these lensing cluster analyses. All of the lensing cluster surveys considered here include sources detected with low significance ($S/N \geq 3.0$), and thus could include a significant number of false positives which would bias the number counts high. However, we believe that the dominate systematic that causes the discrepancies between our number counts and those derived from these lensing cluster surveys is due to the fact that none of these analyses, excepting that of Knudsen et al. (2008), account for the flux boosting effect described in Section 4.3. Knudsen et al. (2008) apply a correction for flux boosting only for the sources detected in their deepest maps ($S_{850\mu\text{m}} \lesssim 2\text{ mJy}$). If we fit for R using only the Knudsen et al. (2008) number counts at $S_{850\mu\text{m}} \leq 2\text{ mJy}$ (row 5 in Table 6), the best-fitting value agrees with our conservative upper limit of $R = 2.5$ within the 1σ errors. However, this effect arises from the steep number counts distribution for SMGs: it is more likely that a source detected in a flux-limited survey is an intrinsically faint source (numerous) boosted high by noise than an intrinsically bright source (rare). Flux boosting is thus prominent in *all* maps, and we believe that the Knudsen et al. (2008) number counts have not been fully corrected for this effect. Still, we note that if SMGs cluster strongly on small angular scales, the apparent bias in lensing cluster surveys may simply be the result of cosmic variance. The total number of lensing clusters analysed is still small, and the errors on the cumulative counts for the lensing cluster surveys are dominated by Poisson noise due to the limited survey areas ($\leq 40\text{ arcmin}^2$ in the source plane) and small sample sizes.

6.4 Comparison with other 1.1 mm surveys

The differential and cumulative number counts from other 1.1 mm blank-field surveys are shown in Figs 7 and 8 for comparison. These

include the AzTEC surveys of GOODS-N (Perera et al. 2008), COSMOS (Scott et al. 2008; Austermann et al. 2009) and SHADES (Austermann et al. 2010) taken on the JCMT. The SHADES data set consists of two separate regions of sky – the Lockman Hole-East and the Subaru/*XMM-Newton* Deep Field – and covers a total area of 0.67 deg^2 , making it the largest blank-field survey at 1.1 mm published to date. The 95 per cent confidence interval from a fluctuation analysis of the 1.1 mm Bolocam Lockman Hole survey ($1\sigma = 1.4\text{ mJy}$; Laurent et al. 2005; Maloney et al. 2005), where a single power-law model is assumed, is also shown. In Fig. 8, we show the 1.2 mm cumulative number counts derived from the MAMBO surveys of the Lockman Hole and ELAIS N2 fields (Greve et al. 2004).

The best-fitting single power-law model for the 1.1 mm number counts determined from our fluctuation analysis of the AzTEC/GOODS-S map is in good agreement with that determined from the fluctuation analysis of the Bolocam Lockman Hole survey. However, given that the AzTEC/GOODS-S survey is 2–3 times deeper than the Bolocam LH survey with roughly the same area, we are sceptical of the extremely tight confidence intervals reported by Maloney et al. (2005). The single power-law model has been shown to poorly describe the SMG number counts over a wide range in flux density (e.g. Coppin et al. 2006; Scott et al. 2006; Austermann et al. 2010), with the best-fitting slope likely a compromise between the steeper slope at high flux densities and the shallower slope at low fluxes. The resulting model inflexibility is a likely contributor to the underrepresented errors. While the 1.2 mm number counts from Greve et al. (2004) are broadly consistent with the number counts determined from several AzTEC surveys, we note that this analysis does not fully account for the substantial flux boosting effect on the measured source flux densities.

The general agreement among the 1.1 mm blank-field surveys taken with AzTEC is quite good. The identical, well-tested methods used to reduce and analyse these data sets minimize any systematic differences among the results, so this agreement is naturally expected. For the same reason, any subtle differences seen among these surveys are significant. The number counts from the GOODS-S field are consistent with those from the SHADES survey, which currently provides the tightest constraints on the blank-field 1.1 mm number counts over this flux density range. In contrast, the GOODS-N and COSMOS fields appear somewhat overdense by as much as ~ 50 per cent; we see a significant amount of variation in the number counts on $\sim 100\text{ arcmin}^2$ scales. Since galaxies are organized within structures with significant power on comoving scales $\lesssim 100\text{ Mpc}$ (e.g. Springel et al. 2005), variations in the observed number counts are expected for surveys covering $\lesssim 1\text{ deg}^2$.⁴ A quantitative study of the degree of field-to-field variations observed among all blank-field surveys taken with AzTEC on the JCMT and the ASTE will be presented in a future paper (Scott et al., in preparation).

7 COMPARISON WITH THE LABOCA SURVEY OF THE ECDF-S AT 870 μm

A 0.35 deg^2 region that covers the ECDF-S has been imaged at 870 μm with the LABOCA camera on the APEX telescope (Weiß et al. 2009). The LABOCA ECDF-S Submillimetre Survey (LESS) therefore encompasses the entire region mapped by AzTEC at 1.1 mm and reaches a uniform complementary depth of

⁴ $10' = 5.1\text{ Mpc}$ at $z = 2$ in the concordance model of cosmology.

$1\sigma \approx 1.2 \text{ mJy beam}^{-1}$ at $870 \mu\text{m}$. A detailed comparison involving a joint analysis of the two data sets by a combined group from the two surveys is in progress. In this section, we carry out a preliminary comparison of the sub-mm/mm sources and their number counts in order to check for general consistency between these two data sets.

7.1 AzTEC- and LABOCA-detected sources

In Fig. 9, we show the AzTEC 1.1 mm map with the locations of the AzTEC and LABOCA sources. Of the 126 $870 \mu\text{m}$ source candidates detected with $S/N \geq 3.7$ in the LABOCA catalogue (where five false detections are expected), 20 are located within the 50 per cent uniform coverage region of the AzTEC survey. Using a conservative search radius of 30 arcsec (roughly the FWHM resolution of both surveys), we find that 16/20 LABOCA sources are also detected by AzTEC with $S/N \geq 3.5$. Turning this around, 16/41 AzTEC source candidates are detected by LABOCA with $S/N \geq 3.7$. Additionally, all seven of the $S/N \geq 4.5$ AzTEC source candidates located in the extended (i.e. noisier) region of our map are also detected by LABOCA. AzTEC sources with probable $870 \mu\text{m}$ counterparts from Weiß et al. (2009) are noted in Table 1. AzTEC/GS42 and AzTEC/GS43 each coincides with pairs of LABOCA source candidates separated by < 25 arcsec.

Verifying that the degree of overlap between the two source catalogues is consistent with expectations given the completeness of the surveys is not straightforward given the different methods used to extract sources and quantify the map properties for these different surveys. This will be fully addressed in the joint analysis of both data sets. The primary reasons for the lower fraction of AzTEC sources detected by LABOCA are (1) we use a lower S/N threshold than that used for the LABOCA catalogue and (2) the AzTEC map (scaled in flux density assuming $S_{870 \mu\text{m}}/S_{1.1 \text{ mm}} = 2$) is slightly deeper than the LABOCA map. We stress that we do not expect a large fraction of the AzTEC-identified sources to be random noise peaks.

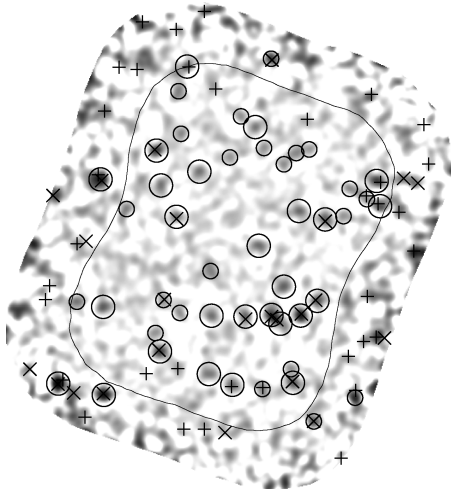


Figure 9. The AzTEC/GOODS-S flux density map with the positions of AzTEC and LABOCA source candidates indicated. To distinguish between higher- and lower-significance detections, AzTEC sources detected with $S/N \geq 5.0$ are shown by larger circles, while AzTEC sources detected with $S/N < 5.0$ are shown by smaller circles. LABOCA sources detected with $S/N \geq 5.0$ are indicated by crosses, and LABOCA sources detected with $S/N < 5.0$ are indicated by pluses. The contour outlines the 50 per cent uniform coverage region for the AzTEC/GOODS-S survey.

For 21 sources that are listed in both the LABOCA and AzTEC source catalogues, we compute the $870 \mu\text{m}$ to 1.1 mm flux ratio from the deboosted fluxes (we exclude AzTEC/GS42 and AzTEC/GS43 since each is resolved into two LABOCA sources). The mean flux ratio of this sample is $(S_{870 \mu\text{m}}/S_{1.1 \text{ mm}}) = 2.0 \pm 0.6$. Scaling from $870 \mu\text{m}$ to $850 \mu\text{m}$ assuming a spectral index of $\beta = 1.5$ ($S_\nu \propto \nu^{3.5}$), this corresponds to a flux ratio of $(S_{850 \mu\text{m}}/S_{1.1 \text{ mm}}) = 2.1 \pm 0.6$, consistent with previous estimates from Chapin et al. (2009) and from scaling the $850 \mu\text{m}$ and 1.1 mm number counts derived from other surveys (Perera et al. 2008; Austermann et al. 2010). The observed flux ratios of the AzTEC/LABOCA sources are also consistent with expectations for cold ($T_d = 25\text{--}50 \text{ K}$), dusty galaxies at $z \sim 1\text{--}3$. From this direct source comparison, there appears to be no significant systematic differences (e.g. calibration) between these two data sets. Note, however, that since this estimate of the average $870 \mu\text{m}$ to 1.1 mm flux ratio is calculated solely from LABOCA-detected AzTEC sources, it may be biased high if there is a significant population of very high-redshift (or very cold) dusty galaxies that are faint at $870 \mu\text{m}$.

7.2 $870 \mu\text{m}$ and 1.1 mm number counts

Weiß et al. (2009) carry out a $P(d)$ analysis on this data set assuming four different parameterized models. Since all of the best-fitting models give nearly the same results for the magnitude and shape of the number counts, we consider only the Schechter function parametrization for the purposes of comparing the Weiß et al. (2009) results to ours. We note that since the LESS field covers a much larger area than the AzTEC/GOODS-S field, we might expect some differences in the source counts from these two data sets owing to cosmic variance. The best-fitting Schechter function to the LABOCA/ECDF-S data, scaled to 1.1 mm assuming a flux ratio of $(S_{870 \mu\text{m}}/S_{1.1 \text{ mm}}) = 2.0$ (as measured from the direct source comparison in Section 7.1), is shown in Figs 7 and 8. The number counts from the LABOCA/ECDF-S field derived using the semi-Bayesian method of Coppin et al. (2006) are also included in these figures.⁵

Considering first the results from the fluctuation analyses, we see fairly good agreement between the number counts derived from the LESS and AzTEC/GOODS-S surveys until $S_{1.1 \text{ mm}} \lesssim 1.0 \text{ mJy}$, at which point the LESS number counts rise more steeply. This is largely due to the differences in the analyses: while we fix $\alpha = -2.0$, Weiß et al. (2009) fit for this parameters and find a best-fitting value of $\alpha = -3.7$ (as defined by equation 3). As explained in Section 6.1, we choose to fix α to a value consistent with that measured by previous AzTEC surveys because the GOODS-S data cannot well constrain a three-parameter model. It is likely that the larger LABOCA/ECDF-S survey is able to put useful constraints on α ; however, Weiß et al. (2009) do not include error estimates on their best-fitting parameters, so it is not possible to quantitatively access the degree to which our results are consistent. We note that the value of α determined from the LESS fluctuation analysis is well matched to our best-fitting spectral index from a fluctuation analysis of AzTEC/GOODS-S assuming a single-power law model.

Comparing the Bayesian-extracted number counts from the AzTEC/GOODS-S and LESS surveys, we find that they are in excellent agreement for $S_{1.1 \text{ mm}} \gtrsim 3 \text{ mJy}$; however, at $S_{1.1 \text{ mm}} \lesssim 3 \text{ mJy}$, the Bayesian-extracted number counts from the LABOCA data are significantly lower than our estimates. Weiß et al. (2009) noted

⁵ The Bayesian-derived number counts from LESS shown in Figs 7 and 8 were provided by A. Weiß.

a similar difference between their $P(d)$ and Bayesian-estimated counts and speculated that this discrepancy at low fluxes arises from source blending, since two or more faint sources blended together by the large beam would be counted as a single bright source and result in an underestimate of the number counts of SMGs at low flux densities. For this reason, Weiß et al. (2009) place more confidence in their $P(d)$ results. However, we have shown through simulations that the $P(d)$ method can be highly model dependent. As described in Section 6.2, our simulations indicate that we largely account for confusion effects in our Bayesian-extracted number counts and that they are not significantly biased low (high) at faint (bright) flux densities due to source blending in the catalogue. We thus place more confidence in our Bayesian-derived results. Since Weiß et al. (2009) do not present similar tests of either of their number counts extraction algorithms, it is not possible to know the underlying cause of the discrepancies between their two different methods.

Weiß et al. (2009) report that the $870\ \mu\text{m}$ number counts from LESS are a factor of ~ 2 lower than those from the $850\ \mu\text{m}$ SCUBA/SHADES survey at flux densities $S_{870\ \mu\text{m}} > 3\ \text{mJy}$ and note that this seems consistent with the reported under-densities in the ECDF-S of other high-redshift galaxy populations. In contrast, we find that the $1.1\ \text{mm}$ number counts from AzTEC/GOODS-S are completely consistent with those measured in the AzTEC/SHADES survey. Since sub-mm/mm-selected galaxies potentially span a much larger volume of space than high-redshift galaxies selected by other photometric criterion, we would not necessarily expect to see an underdensity of SMGs in this field. As the number counts derived from the LESS, AzTEC/GOODS-S and AzTEC/SHADES fields are largely consistent assuming a reasonable scaling for the differences in observed wavelength, it seems unlikely that cosmic variance alone can account for the factor of ~ 2 difference in the number counts derived from LESS and SCUBA/SHADES. Considering also the physically unrealistic high flux-scaling factor required to bring the SCUBA/SHADES and AzTEC/SHADES number counts into agreement ($R = 2.5$), one possibility is that this discrepancy arises from systematics in the SCUBA/SHADES calibration and/or number counts extraction as discussed by Austermann et al. (2010); this will be properly addressed via a source to source comparison of AzTEC and SCUBA sources in the SHADES fields (Negrello et al., in prep.).

8 CONTRIBUTION OF DIFFERENT GALAXY POPULATIONS TO THE COSMIC INFRARED BACKGROUND AT 1.1 MM

The FIRAS instrument on the *Cosmic Background Explorer* (COBE) detected a uniform cosmic infrared background (CIRB), which peaks around $200\ \mu\text{m}$ (Puget et al. 1996; Fixsen et al. 1998). The total power in the infrared background is 1–2.6 times larger than that of the optical/UV background, suggesting that the majority of radiation produced by young stars and AGN over the cosmic history is absorbed and reprocessed by dust (e.g. Gispert, Lagache & Puget 2000). If dusty starburst galaxies in the early universe account for much of the CIRB, the AzTEC/GOODS-S survey is expected to resolve a significant fraction of the CIRB.

Summing the deboosted $1.1\ \text{mm}$ flux densities of the 41 $S/N \geq 3.5$ source candidates in the AzTEC/GOODS-S map, we measure an integrated flux of $1.5 \pm 0.2\ \text{Jy deg}^{-2}$ over the $0.075\ \text{deg}^2$ field. Comparing this to the total energy density in the CIRB at $1.1\ \text{mm}$ of $18\text{--}24\ \text{Jy deg}^{-2}$ (Puget et al. 1996; Fixsen et al. 1998), we have resolved only 6–8 per cent of the CIRB into individual galaxies. However, if we instead integrate the best-fitting Schechter function

model to the GOODS-S number counts down to $0\ \text{mJy}$, we estimate that we have statistically detected $6.3\ \text{Jy deg}^{-2}$, or 26–35 per cent, of the CIRB at $1.1\ \text{mm}$ with our AzTEC/GOODS-S survey. The fact that we do not resolve 100 per cent of the CIRB through this integration implies that simply extrapolating the Schechter function model that best fits our data at $0.5\ \text{mJy} < S_{1.1\ \text{mm}} < 6.5\ \text{mJy}$ to much lower flux densities significantly underestimates the number counts at the faint end. On the other hand, integrating the best-fitting single power-law model to the GOODS-S number counts down to $S_{1.1\ \text{mm}} = 0.04\ \text{mJy}$ would account for 100 per cent of the CIRB at $1.1\ \text{mm}$; however, we know this model to be a poor description of the number counts of bright SMGs as well as physically unfeasible since it approaches infinity as $S_{1.1\ \text{mm}}$ goes to zero.

We next use a stacking analysis to estimate the fraction of the CIRB at $1.1\ \text{mm}$ that is resolved by the entire $1.4\ \text{GHz}$ radio population. Stacking at the locations of $N = 222$ radio sources in this field (up slightly from 219 radio sources used in the stacking analysis in Section 4.2, since we have shifted our AzTEC map to correct the astrometry), we calculate an average $1.1\ \text{mm}$ flux density of $S_{1.1\ \text{mm}, \text{radio}} = 660 \pm 78\ \mu\text{Jy}$. Assuming that each of the radio sources distributed over an area of $A = 0.075\ \text{deg}^2$ has a flux density of $S_{1.1\ \text{mm}, \text{radio}}$, the radio population has a total integrated flux of $N \cdot S_{1.1\ \text{mm}, \text{radio}} / A = 1.9\ \text{Jy deg}^{-2}$, or 8–11 per cent of the CIRB at $1.1\ \text{mm}$.

We finally estimate the contribution to the CIRB at $1.1\ \text{mm}$ from MIPS $24\ \mu\text{m}$ -selected sources using a similar stacking analysis. The average flux density from 1185 $24\ \mu\text{m}$ sources distributed over a $0.068\ \text{deg}^2$ area is $S_{1.1\ \text{mm}, 24\ \mu\text{m}} = 290 \pm 26\ \mu\text{Jy}$. The total integrated flux from $24\ \mu\text{m}$ sources at $1.1\ \text{mm}$ is $5.0\ \text{Jy deg}^{-2}$, or 21–28 per cent of the total CIRB. This is similar to the fraction of the CIRB at $850\ \mu\text{m}$ resolved by $24\ \mu\text{m}$ sources (29–37 per cent) in the SCUBA/GOODS-N field (Wang, Cowie & Barger 2006). In contrast, a stacking analysis of $24\ \mu\text{m}$ sources with the BLAST maps of GOODS-S at 250, 350 and $500\ \mu\text{m}$ suggests that the full intensity of the CIRB at these shorter wavelengths is resolved by sources selected at $24\ \mu\text{m}$ (Devlin et al. 2009; Marsden et al. 2009). This possibly demonstrates the existence of a significant population of higher-redshift ($z \gtrsim 3$) dust-obscured galaxies that are (statistically speaking) missed by current $\lambda \lesssim 500\ \mu\text{m}$ surveys, but account for $\approx 2/3$ of the CIRB at longer wavelengths.

9 CONCLUSIONS

We imaged a $270\ \text{arcmin}^2$ field towards the GOODS-S region to a confusion-limited depth of $1\sigma \approx 0.6\ \text{mJy}$ using the AzTEC camera on the ASTE, making this the deepest survey carried out to date at $1.1\ \text{mm}$. We detect 41 SMG candidates with $S/N \geq 3.5$, where roughly two are expected to be false positives arising from noise peaks. This survey is 50 per cent complete at $2.1\ \text{mJy}$ and 95 per cent complete at $3.5\ \text{mJy}$. We have demonstrated that the presence of confusion noise has significant consequences for the properties of the map (and hence the number counts) and must be considered when accessing the survey completeness and expected number of false detections in the source catalogue. From realistic simulations of the SMG population in this field, we estimate that ≈ 26 per cent of the source candidates identified in the AzTEC/GOODS-S map are actually two or more mm-bright galaxies with comparable flux densities blended together due to the low angular resolution of the ASTE beam.

We have used two very different methods to estimate the SMG number counts in this field: a fluctuation analysis where we model the distribution of flux density values in the map and a

semi-Bayesian technique where the number counts are determined from sampling the PFDs from the catalogue of SMGs. We have demonstrated that both methods are able to retrieve the correct number counts distribution from fully simulated data sets. Furthermore, the best-fitting number counts to the GOODS-S field using these different methods are consistent. We note that while our fluctuation analysis can provide good constraints on the number counts, the results depend strongly on the assumed model for the number counts distribution. The depth and large survey area of the AzTEC/GOODS-S map have resulted in the tightest constraints to date on the SMG number counts at $S_{1.1\text{ mm}} = 0.5\text{ mJy}$. Comparing our cumulative number counts to those from several SCUBA lensing cluster surveys at $850\text{ }\mu\text{m}$, the lensing cluster number counts appear to be biased high assuming reasonable values for the flux-scaling from $850\text{ }\mu\text{m}$ to 1.1 mm . These results are consistent with those from large-area blank-field surveys at $850\text{ }\mu\text{m}$ with SCUBA, where the number counts at $S_{850\text{ }\mu\text{m}} \gtrsim 2\text{ mJy}$ from lensing cluster surveys are systematically higher; this is most likely due to flux boosting effects not being fully treated in the number counts extraction from the lensing cluster surveys. We find that the number counts from the AzTEC/GOODS-S field are consistent with those from the 0.67 deg^2 AzTEC/SHADES survey.

Comparing our source catalogue to the $870\text{ }\mu\text{m}$ source list from the LABOCA/ECDF-S survey, 16/20 (80 per cent) LABOCA sources located within the 50 per cent uniform coverage region of the AzTEC/GOODS-S survey are also detected with $S/N \geq 3.5$ in the AzTEC map; however, since the AzTEC map is slightly deeper, only 16/41 (39 per cent) of the AzTEC sources in this field are listed in the LABOCA source catalogue. In contrast to the results from the LABOCA survey, we find no apparent underdensity of SMGs in this field compared to previous sub-mm/mm surveys. For the SMGs that are listed in both catalogues, the mean $870\text{ }\mu\text{m}$ to 1.1 mm flux ratio is 2.0 ± 0.6 and is consistent with our expectations given typical dust SEDs and redshifts for these galaxies. Scaling the $870\text{ }\mu\text{m}$ number counts to 1.1 mm , we find fairly good agreement between the counts for $S_{1.1\text{ mm}} \gtrsim 1.0\text{ mJy}$ ($\gtrsim 3.0\text{ mJy}$) extracted using a fluctuation analysis (Bayesian method) on each of the two surveys. We think that the discrepancies between the AzTEC and LABOCA source counts at fainter flux densities are due to differences in the assumed models used for the fluctuation analyses and additional biases in the Bayesian method for the LABOCA analysis as noted by Weiß et al. (2009). A combined analysis of these two confusion-limited surveys at $870\text{ }\mu\text{m}$ and 1.1 mm is in progress.

We resolve only 6–8 per cent of the CIRB at 1.1 mm into individual mm-bright galaxies. While the $24\text{ }\mu\text{m}$ population can account for the full energy density in the CIRB at $250\text{--}500\text{ }\mu\text{m}$, we estimate that $24\text{ }\mu\text{m}$ sources statistically detect only 21–28 per cent of the CIRB at 1.1 mm , demonstrating that a population of faint dust-obscured galaxies at $z \gtrsim 3$ that are largely missed at shorter wavelengths potentially contribute significantly to the total energy density in the CIRB at 1.1 mm .

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REFERENCES

- Austermann J. E. et al., 2009, MNRAS, 393, 1573
- Austermann J. E. et al., 2010, MNRAS, 401, 160
- Baugh C. M., Lacey C. G., Frenk C. S., Granato G. L., Silva L., Bressan A., Benson A. J., Cole S., 2005, MNRAS, 356, 1191
- Bertoldi F. et al., 2007, ApJS, 172, 132
- Blain A. W., Smail I., Ivison R. J., Kneib J.-P., Frayer D. T., 2002, Phys. Rep., 369, 111
- Blain A. W., Chapman S. C., Smail I., Ivison R., 2004, ApJ, 611, 725
- Borys C., Chapman S., Halpern M., Scott D., 2003, MNRAS, 344, 385
- Chapin E. L. et al., 2009, MNRAS, 398, 1793
- Chapman S. C., Scott D., Borys C., Fahlman G. G., 2002, MNRAS, 330, 92
- Chapman S. C., Blain A. W., Smail I., Ivison R. J., 2005, ApJ, 622, 772
- Coppin K., Halpern M., Scott D., Borys C., Chapman S., 2005, MNRAS, 357, 1022
- Coppin K. et al., 2006, MNRAS, 372, 1621
- Coppin K. et al., 2008, MNRAS, 384, 1597
- Cowie L. L., Barger A. J., Kneib J.-P., 2002, AJ, 123, 2197
- Devlin M. J. et al., 2009, Nat, 458, 737
- Dunne L., Eales S. A., 2001, MNRAS, 327, 697
- Dunne L., Eales S., Edmunds M., Ivison R., Alexander P., Clements D. L., 2000, MNRAS, 315, 115
- Eales S., Lilly S., Webb T., Dunne L., Gear W., Clements D., Yun M., 2000, AJ, 120, 2244
- Eales S., Bertoldi F., Ivison R., Carilli C., Dunne L., Owen F., 2003, MNRAS, 344, 169
- Ezawa H., Kawabe R., Kohno K., Yamamoto S., 2004, in Oschmann J. M., Jr, ed., Proc. SPIE Vol. 5489, Ground-based Telescopes. SPIE, Bellingham, p. 763
- Ezawa H. et al., 2008, in Stepp L. M., Gilmozzi R., eds, Proc. SPIE Vol. 7012, Ground-based and Airborne Telescopes II. SPIE, Bellingham, p. 701208
- Fixsen D. J., Dwek E., Mather J. C., Bennett C. L., Shafer R. A., 1998, ApJ, 508, 123
- Gialalisco M. et al., 2004, ApJ, 600, L93
- Gispert R., Lagache G., Puget J. L., 2000, A&A, 360, 1
- Glenn J. et al., 1998, in Phillips T. G., ed., Proc. SPIE Vol. 3357, Advanced Technology MMW, Radio, and Terahertz Telescopes. SPIE, Bellingham, p. 326
- Granato G. L., De Zotti G., Silva L., Bressan A., Danese L., 2004, ApJ, 600, 580
- Greve T. R. et al., 2005, MNRAS, 359, 1165

- Greve T. R., Ivison R. J., Bertoldi F., Stevens J. A., Dunlop J. S., Lutz D., Carilli C. L., 2004, *MNRAS*, 354, 779
- Greve T. R., Pope A., Scott D., Ivison R. J., Borys C., Conselice C. J., Bertoldi F., 2008, *MNRAS*, 389, 1489
- Griffin M. J., Orton G. S., 1993, *Icarus*, 105, 537
- Haig D. J. et al., 2004, in Zmuidzinas J., Holland W. S., Withington S., eds, *Proc. SPIE Vol. 5498, Millimeter and Submillimeter Detectors for Astronomy II*. SPIE, Bellingham, p. 78
- Hogg D. W., 2001, *AJ*, 121, 1207
- Hogg D. W., Turner E. L., 1998, *PASP*, 110, 727
- Holland W. S. et al., 1999, *MNRAS*, 303, 659
- Hughes D. H. et al., 1998, *Nat*, 394, 241
- Ivison R. J. et al., 2005, *MNRAS*, 364, 1025
- Ivison R. J. et al., 2007, *MNRAS*, 380, 199
- Kamazaki T. et al., 2005, in Shopbell P., Britton M., Ebert R., eds, *ASP Conf. Ser. Vol. 347, Astronomical Data Analysis Software and Systems XIV*. Astron. Soc. Pac., San Francisco, p. 533
- Kellermann K. I., Fomalont E. B., Mainieri V., Padovani P., Rosati P., Shaver P., Tozzi P., Miller N., 2008, *ApJS*, 179, 71
- Knudsen K. K. et al., 2006, *MNRAS*, 368, 487
- Knudsen K. K., van der Werf P. P., Kneib J.-P., 2008, *MNRAS*, 384, 1611
- Kovács A., Chapman S. C., Dowell C. D., Blain A. W., Ivison R. J., Smail I., Phillips T. G., 2006, *ApJ*, 650, 592
- Kreysa E. et al., 1998, in Phillips T. G., ed., *Proc. SPIE Vol. 3357, Advanced Technology MMW, Radio, and Terahertz Telescopes*. SPIE, Bellingham, p. 319
- Laurent G. T. et al., 2005, *ApJ*, 623, 742
- Luo B. et al., 2008, *ApJS*, 179, 19
- Maloney P. R. et al., 2005, *ApJ*, 635, 1044
- Marsden G. et al., 2009, *ApJ*, 707, 1729
- Miller N. A., Fomalont E. B., Kellermann K. I., Mainieri V., Norman C., Padovani P., Rosati P., Tozzi P., 2008, *ApJS*, 179, 114
- Perera T. A. et al., 2008, *MNRAS*, 391, 1227
- Pope A. et al., 2006, *MNRAS*, 370, 1185
- Popesso P. et al., 2009, *A&A*, 494, 443
- Puget J.-L., Abergel A., Bernard J.-P., Boulanger F., Burton W. B., Desert F.-X., Hartmann D., 1996, *A&A*, 308, L5
- Retzlaff J., Rosati P., Dickinson M., Vandame B., Rite C., Nonino M., Cesarsky C., the GOODS Team, 2010, *A&A*, 511, 50
- Rowan-Robinson M., 2009, *MNRAS*, 394, 117
- Sanders D. B., Mirabel I. F., 1996, *ARA&A*, 34, 749
- Scott S. E. et al., 2002, *MNRAS*, 331, 817
- Scott S. E., Dunlop J. S., Serjeant S., 2006, *MNRAS*, 370, 1057
- Scott K. S. et al., 2008, *MNRAS*, 385, 2225
- Silva L., De Zotti G., Granato G. L., Maiolino R., Danese L., 2005, *MNRAS*, 357, 1295
- Smail I., Ivison R. J., Blain A. W., Kneib J.-P., 1998, *ApJ*, 507, L21
- Smail I., Ivison R. J., Blain A. W., Kneib J.-P., 2002, *MNRAS*, 331, 495
- Springel M. et al., 2005, *Nat*, 435, 629
- Stevens J. A. et al., 2003, *Nat*, 425, 264
- Takeuchi T. T., Ishii T. T., 2004, *ApJ*, 604, 40
- Tamura Y. et al., 2009, *Nat*, 459, 61
- Vanzella E. et al., 2005, *A&A*, 434, 53
- Vanzella E. et al., 2006, *A&A*, 454, 423
- Vanzella E. et al., 2008, *A&A*, 478, 83
- Wang W.-H., Cowie L. L., Barger A. J., 2006, *ApJ*, 647, 74
- Webb T. M. et al., 2003, *ApJ*, 587, 41
- Weiß A. et al., 2009, *ApJ*, 707, 1201
- Wilson G. W. et al., 2008, *MNRAS*, 386, 807
- Younger J. D. et al., 2007, *ApJ*, 671, 1531
- Younger J. D. et al., 2009, *ApJ*, 704, 803

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